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Messineo, Maria; Clark, J. Simon; Figer, Donald F.; Kudritzki, Rolf-Peter; Najarro, Francisco; Rich, R. Michael; Menten, Karl M.; Ivanov, Valentin D.; Valenti, Elena; Trombley, Christine; Chen, C.-H. Rosie and Davies, Ben (2015). Massive Stars In The W33 Giant Molecular Complex. *The Astrophysical Journal*, 805(2) p. 110.

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<http://dx.doi.org/doi:10.1088/0004-637X/805/2/110>

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MASSIVE STARS IN THE W33 GIANT MOLECULAR COMPLEX

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Received 2015 January 30; accepted 2015 April 1; published 2015 May 27

ABSTRACT

Rich in H II regions, giant molecular clouds are natural laboratories to study massive stars and sequential star formation. The Galactic star-forming complex W33 is located at $l = \sim 12^\circ 8$ and at a distance of 2.4 kpc and has a size of ~ 10 pc and a total mass of $\approx (0.8\text{--}8.0) \times 10^5 M_\odot$. The integrated radio and IR luminosity of W33—when combined with the direct detection of methanol masers, the protostellar object W33A, and the protocluster embedded within the radio source W33 main—mark the region as a site of vigorous ongoing star formation. In order to assess the long-term star formation history, we performed an infrared spectroscopic search for massive stars, detecting for the first time 14 early-type stars, including one WN6 star and four O4–7 stars. The distribution of spectral types suggests that this population formed during the past $\sim 2\text{--}4$ Myr, while the absence of red supergiants precludes extensive star formation at ages 6–30 Myr. This activity appears distributed throughout the region and does not appear to have yielded the dense stellar clusters that characterize other star-forming complexes such as Carina and G305. Instead, we anticipate that W33 will eventually evolve into a loose stellar aggregate, with Cyg OB2 serving as a useful, albeit richer and more massive, comparator. Given recent distance estimates, and despite a remarkably similar stellar population, the rich cluster Cl 1813–178 located on the northwest edge of W33 does not appear to be physically associated with W33.

Key words: infrared: stars – stars: evolution

1. INTRODUCTION

Massive stars enrich the galactic interstellar medium via the feedback of radiative and mechanical energy, the deposition of chemically processed gas via their strong winds, and, latterly, solid state material during the post-main-sequence (post-MS) phase. Because of their luminosities, individual massive stars can be detected and resolved in external galaxies, providing direct measures of distances and spatially resolved metallicity gradients (e.g., Kudritzki et al. 2014). At and beyond the end of their lives they power a wide variety of highly energetic transient phenomena—first during their deaths in supernovae (SNe) or gamma-ray bursts and subsequently by accretion onto their stellar corpses in X-ray binaries (e.g., Güdel & Nazé 2009; Eldridge et al. 2013).

Considerable uncertainty remains regarding the mechanism(s) of formation of massive stars, although it is strongly suspected that this process is hierarchical: massive stars are found in apparently isolated young massive stellar clusters (e.g., the Arches and Quintuplet clusters; Figer et al. 2002), in loose associations (e.g., Cyg OB2; Negueruela et al. 2008; Wright et al. 2014, and references therein), and in large molecular complexes (e.g., 30 Doradus and G305; Walborn & Blades 1997; Clark & Porter 2004).

Massive stars are very often part of binary systems (typically, a fraction of 91% of OB stars are found to have

companions; Sana et al. 2014). A population of apparently isolated massive stars also exists, although it is not clear whether these have genuinely formed in isolation (e.g., Bestenlehner et al. 2011) or, instead, were lost from a natal aggregate due to dynamical or SN-driven ejection (runaway stars; Povich et al. 2008; Oh et al. 2014).

Because of the high and variable interstellar extinction and uncertain distances of stars within the Galactic disk, it has long been suspected that our census of massive star-forming regions is incomplete. Fortunately, the plethora of modern infrared and radio surveys—e.g., MAGPIS, GLIMPSE, WISE, MSX, 2MASS, UKIDSS, and VVV,¹²—allow us to identify both the natal giant molecular clouds (GMCs) and the stars that form within them. Subsequent analysis of the physical properties of GMCs and associated stellar population(s)—in terms of the mass function of pre-stellar clumps/cores (proto-stars) and already formed massive stars, and their temporal and

¹² MAGPIS stands for the Multi-Array Galactic Plane Imaging Survey (White et al. 2005; Helfand et al. 2006), 2MASS for Two Micron All Sky Survey (Skrutskie et al. 2006), DENIS for Deep Near Infrared Survey of the Southern Sky (Epstein et al. 1994), UKIDSS for UKIRT Infrared Deep Sky Survey (Lucas et al. 2008), VVV for the VISTA Variables in the Via Lactea survey (Soto et al. 2013), MSX for Midcourse Space Experiment (Egan et al. 2003; Price et al. 2001), GLIMPSE for Galactic Legacy Infrared Mid-Plane Survey Extraordinaire (Churchwell et al. 2009), and WISE for Wide-field Infrared Survey Explorer (Wright et al. 2010).

spatial distributions—enable constraints to be placed on the mode of star formation that occurred in the region in question (e.g., Messineo et al. 2014a).

One such massive star-forming region is the W33 complex, located in the Galactic plane at longitude $l = \sim 12^\circ 8'$; a parallactic distance of $2.40^{+0.17}_{-0.15}$ kpc was determined from observations of water masers, which suggests a location in the Scutum spiral arm (Immer et al. 2013). Subtending $15'$ (~ 10 pc), it comprises a number of distinct molecular and/or dusty condensations (see Immer et al. 2014 for a census), with an integrated IR luminosity of $\sim 8 \times 10^5 L_\odot$ and a total mass of $\sim (0.8\text{--}8) \times 10^5 M_\odot$ (Immer et al. 2013). Radio observations of one component—W33 Main—revealed the presence of an obscured (proto-)cluster apparently comprising a number of stars with spectral types ranging from O7.5 to B1.5 (Haschick & Ho 1983). The presence of ongoing massive star formation is also signposted by the presence of OH, H₂O, and CH₃OH masers (Immer et al. 2013) and the direct identification of a bipolar outflow and massive dusty torus associated with the young stellar object W33A (Davies et al. 2010).

Independently of these studies, Messineo et al. (2008, 2011) serendipitously identified a hitherto-overlooked young massive cluster—CI 1813–178—in the vicinity of W33. Analysis of the post-MS content of the cluster suggested a mass $\gtrsim 10^4 M_\odot$, making it among the most massive aggregates in the Galaxy (Clark et al. 2013). Given the unusual mix of spectral types present, Messineo et al. (2011) quoted 4–4.5 Myr but highlighted that several cluster members had low luminosities for that age; stellar luminosities would appear to demonstrate some degree of non-coevality. Intriguingly, CI 1813–178 is found in the vicinity of the pulsar wind nebula HESS J1813–178 (Helfand et al. 2007; Messineo et al. 2008). With a spin-down measurement of 44.7 ms and a spin-down luminosity of $\dot{E} \sim 5.6 \times 10^{37}$ erg s^{−1}, PSR J1813–1749 is one of the youngest and most energetic pulsars in the Galaxy (Halpern et al. 2012). While the energetic young pulsar potentially lies beyond both regions (>4.8 kpc; Halpern et al. 2012), evidently this line of sight samples numerous regions of massive star formation.

Here, we present a near-infrared spectroscopic survey of bright stars in selected regions of W33 (cl1, cl2, and Mercer1; Messineo et al. 2011) and in the nearby GLIMPSE bubble N10 (e.g., Churchwell et al. 2006). We aim to determine the massive stellar content of W33, its star formation history, and, hence, relation to the nearby cluster CI 1813–178 and pulsar PSR J1813–1749. In Section 2 we present the spectroscopic observations, and in Section 3 the available infrared photometry. In Sections 4 and 5 we describe the spectral features and stellar properties. In Sections 6, 7, and 8 we briefly describe the spatial and temporal distributions of the detected stars, before summarizing our findings in Section 9.

2. TARGETS AND SPECTROSCOPIC OBSERVATIONS

Spectroscopic targets with 2MASS K_s from 6 to 11 mag and $H-K_s > 0.5$ mag were selected from the regions listed in Messineo et al. (2011) that exhibited overdensities of bright stars or nebulae in 2MASS and GLIMPSE images. The Mercer1 region,¹³ to the west of the radio source W33 Main, appears as a sparse aggregate of bright stars with a pronounced

arc of IR and radio emission to the south that is suggestive of a wind-blown structure (Figure 1, top panel). The cl1 region, immediately to the east of the embedded proto-cluster in W33 Main, contains an isolated bright star surrounded by an arc of emission at IR and radio wavelengths (Figure 1, middle panel). The cl2 region, coincident with W33 B1 (e.g., Immer et al. 2014), contains a small cluster of stars associated with diffuse IR and radio emission (Figure 1, bottom panel). These targets were supplemented with a few isolated bright targets selected on the basis of their GLIMPSE colors—e.g., indicative of free-free emitters (e.g., Hadfield et al. 2007; Messineo et al. 2012)—such as star #1 and the candidate luminous blue variable (LBV) #15 from CI 1813–178. Additionally, we observed some stars to the north of W33, in the cluster BDS2003-115 embedded in the GLIMPSE mid-IR bubble N10 (Figure 2).

Data were acquired with the Spectrograph for Integral Field Observations in the Near Infrared (SINFONI; Eisenhauer et al. 2003) on the Yepun Very Large Telescope, under the ESO programs 087.D-0265(A) and 089.D-0790(A). The K -grating was used along with the $0''.25$ pixel^{−1} scale to yield a resolving power of ~ 4500 . Integration times per exposure ranged from 30 to 300 s. Typically, each observation consisted of four exposures, two on target and two on sky. Data-cubes were generated with the version 3.9.0 of the ESO SINFONI pipeline (Schreiber et al. 2004; Modigliani et al. 2007), using flat-fields, bad-pixel masks, distortion maps, and arcs. From each cube, stellar traces with signal-to-noise ratio larger than 30–40 were analyzed. Corrections for instrumental and atmospheric responses were accomplished with standard stars of B types; stellar Br _{γ} and He I lines at $2.112 \mu\text{m}$ were removed from the standard spectra with a linear interpolation. A sky subtraction was performed to eliminate possible nebular lines and residuals from OH subtraction. A total of 86 cubes were observed, and 94 stars were extracted.

3. AVAILABLE PHOTOMETRIC DATA

We searched for counterparts of the spectroscopically observed stars in the 2MASS Catalog of Point Sources (Skrutskie et al. 2006), in the UKIDSS catalog (Lucas et al. 2008), and in the DENIS catalog (DENIS Consortium 2005) by taking the closest match within $1''$. Mid-infrared data were retrieved from the MSX (Egan et al. 2003) with a search radius of $5''$, from the GLIMPSE (Churchwell et al. 2009) and the WISE (Wright et al. 2010) with a search radius of $2''$.

For 21% of the sources, R -band counterparts with magnitudes from 12.38 to 19.80 mag were found in the Naval Observatory Merged Astrometric Dataset (NOMAD; Zacharias et al. 2004).

Photometric measurements are listed in the Appendix. For the observed targets, there is no additional information from the SIMBAD database.

3.1. UKIDSS Photometry

For stars fainter than $K_s \sim 10.5$ mag, we used UKIDSS photometry (Lucas et al. 2008). JHK magnitudes are available from the UKIDSS data release number 6 (DR6; Lucas et al. 2008); however, for four fields (Mercer1, cl1, cl2, N10) we generated photometric catalogs of point sources with the psf-fitting algorithm DAOPHOT (Stetson 1987) and the

¹³ Messineo et al. (2011) define this as a region with a radius of $2/3$ that includes the candidate stellar cluster n. 1 of Mercer et al. (2005).

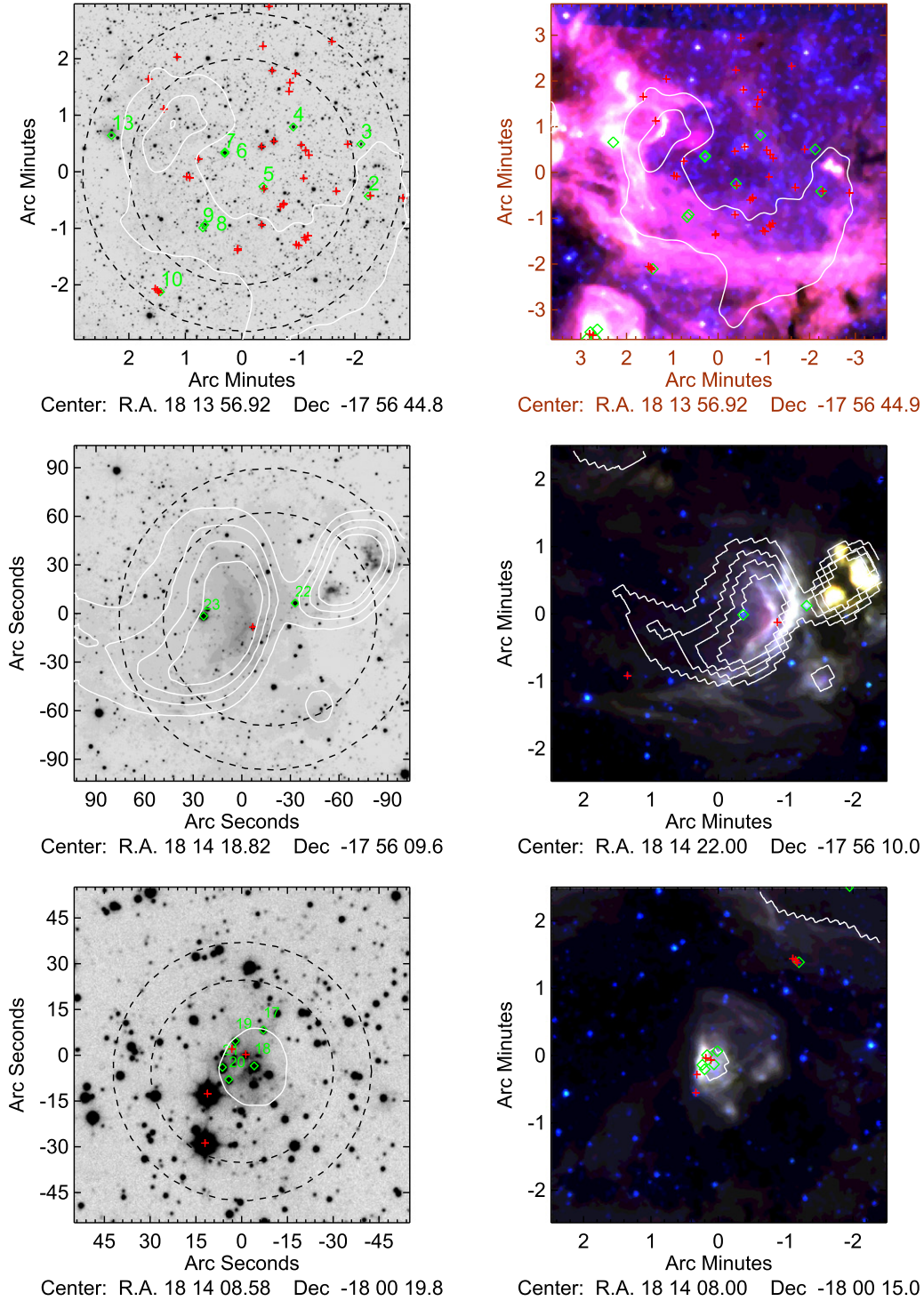


Figure 1. Images of the stellar clusters Mercer1 (top row), c11 (middle row), and c12 (bottom row) in W33 (Messineo et al. 2011). Left side: detected stars are displayed on a UKIDSS *K*-band image (for c12 the *J* image is used). Diamonds and crosses indicate the positions of detected early types and late types, respectively. Contours of the 90 cm radio-continuum emission from 100 up to 400 mJy beam^{-1} (with a step of 100 mJy beam^{-1}) are shown in white. The two dashed circles mark the regions used for the CMDs in Figure 8. By assuming a distance of 2.4 kpc, 1' corresponds to 0.70 pc. Right side: contours of the 90 cm radio continuum emission superimposed on a GLIMPSE composite image (3.6 μm in blue, 5.8 μm in green, and 8.0 μm in red).

leavestack frames provided by UKIDSS (Lucas et al. 2008). A detection threshold of 4σ was used. From overlapping fields, independently calibrated with 2MASS datapoints, an absolute photometric error of 0.05 mag was estimated. Comparison of the UKIDSS pipeline and the psf-fitting catalogs in the Mercer1 field yielded an absolute uncertainty of 0.07 mag. The resultant magnitudes are given in the Appendix.

4. ANALYSIS

4.1. Spectral Classification

We detected a total of 23 new early types, which comprise 1 Wolf-Rayet, 13 stars of spectral type O or B, and 9 stars with indeterminate, but early, spectral types (OBAF), and 70 late-type stars (see Tables 1, 2, and 3).

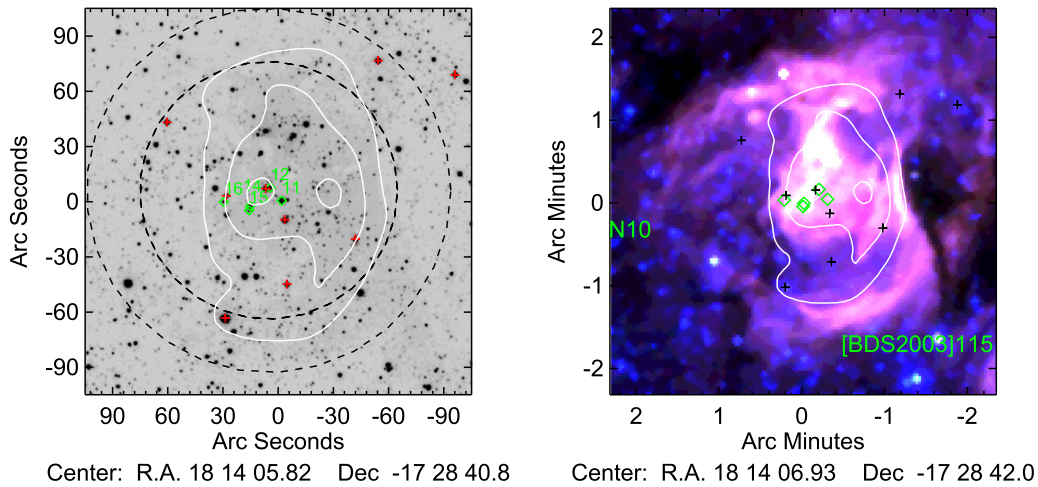


Figure 2. Images of the bubble N10. Left side: detected stars are displayed on a UKIDSS *K*-band image. Diamonds and crosses indicate the positions of detected early types and late types, respectively. The white contours show the 90 cm radio continuum emission of the bubble N10 from MAGPIS at 100, 200, and 300 mJy beam⁻¹. By assuming a distance of 4.29 kpc, 1' corresponds to 1.25 pc. Right side: 90 cm contours are superimposed on a GLIMPSE composite image (3.6 μ m in blue, 5.8 μ m in green, and 8.0 μ m in red).

4.1.1. Early-type Stars

The spectra of the newly discovered early-type stars are characterized by hydrogen and helium lines, as well as transitions of heavier elements, such as C, N, O, and Fe. Their spectral classification was accomplished by comparison to the near-infrared atlases of Hanson et al. (1996), Hanson et al. (2005), and Figer et al. (1997).

Star #1 is a newly discovered Wolf-Rayet star; its spectrum is characterized by strong and broad emission lines: He I λ 2.1117 μ m, He II λ 2.1636 μ m, and He II λ 2.1891 μ m. The equivalent width ratios between the 2.1891 μ m line and the two lines at 2.1636 μ m and 2.112 μ m were estimated to be 1.8 ± 0.3 and 4.3 ± 0.7 , respectively,¹⁴ by using multiple Gaussian fits (Figer et al. 1997). This spectrum resembles that of WR134, a WN6b star (Figer et al. 1997); the suffix b indicates broad emission lines (e.g., Hadfield et al. 2007); derived ratios are almost identical to those calculated for WR134; thereby, star #1 is a WN6b.

The spectra of stars #7, #8, #11, and #23 are characterized by emission in the C IV 2.0705 μ m and 2.0796 μ m and O III/N III at ≈ 2.115 μ m lines, He I line at 2.059 μ m and He II 2.189 μ m in absorption, and Br γ mostly in absorption (see Table 2 and Figure 3). This combination of features is characteristic of stars of mid- to late O spectral type; unfortunately, for this temperature range the assignment of luminosity classes from *K*-band spectra alone is somewhat problematic. Nevertheless, in this regard we note the strong morphological resemblance of these objects to the O4–6 I stars within the Arches cluster (Martins et al. 2008). The spectra of stars #8, #11, and #23 display the Si IV emission line at 2.428 μ m. We denote O-type stars with O III/N III and Si IV in emission by Of_K⁺, as described in Messineo et al. (2014a), although we caution that this does not necessarily indicate a super/hyper-giant classification (de Jager 1988). Star #8 has Br γ filled in, indicating I+ nature. The

¹⁴ Errors on the ratios are calculated by propagating the errors on the EWs; for each line, errors on the EWs are obtained with the formula of Vollmann & Eversberg (2006). For the Br γ and He II line at 2.189 μ m, EWs are calculated after having subtracted from the observed spectrum the Gaussian fits to the contaminating lines.

Table 1
Summary of Observations of Early-type Stars

ID	R.A.[J2000] (hh mm ss)	Decl.[J2000] (deg mm ss)	Spec.	Period	Date-Obs (yyyy mm dd)
M15 ^a	18 13 20.99	-17 49 46.9	cLBV	P89	2012 Jun 21
1	18 13 34.81	-18 05 41.5	WN6	P89	2012 Jun 21
2	18 13 47.53	-17 57 10.7	OBAF	P89	2012 Jun 27
3	18 13 48.07	-17 56 15.6	B0-5	P87	2011 Aug 19
4	18 13 53.10	-17 55 57.3	B0-5	P87	2011 May 29
5	18 13 55.36	-17 57 00.7	OBAF	P87	2011 May 29
6	18 13 58.19	-17 56 23.9	B0-5	P87	2011 May 29
7	18 13 58.20	-17 56 25.4	O4-6	P87	2011 May 29
8	18 13 59.69	-17 57 41.2	O4-6	P87	2011 Aug 19
9	18 13 59.83	-17 57 44.4	OBAF	P87	2011 Aug 20
10	18 14 03.00	-17 58 52.3	B0-5	P87	2011 Aug 20
11	18 14 05.69	-17 28 40.3	O4-6	P87	2011 Jun 24
12	18 14 06.12	-17 28 33.2	OBAF	P87	2011 Jun 24
13	18 14 06.63	-17 56 06.3	B0-5	P87	2011 Aug 26
14	18 14 06.89	-17 28 43.2	OBe	P89	2012 Sep 15
15	18 14 06.93	-17 28 45.5	OBAF	P89	2012 Sep 15
16	18 14 07.89	-17 28 40.9	OBAF	P89	2012 Sep 15
17	18 14 08.09	-18 00 11.6	B0-5	P89	2012 Jun 26
18	18 14 08.30	-18 00 23.3	B0-5	P87	2011 May 29
19	18 14 08.74	-18 00 15.2	OBAF	P89	2012 Jun 26
20	18 14 08.89	-18 00 27.8	OBAF	P89	2012 Sep 15
21	18 14 09.03	-18 00 23.8	OBAF	P89	2012 Sep 15
22	18 14 16.52	-17 56 03.2	Oe	P87	2011 May 29
23	18 14 20.48	-17 56 11.2	O6-7	P87	2011 May 29

Note.

^a Star [MFD2011] 15 was discovered by Messineo et al. (2011).

spectrum of star #23 has the He I line at 2.112 μ m in absorption and most likely has a later sub-type.

The spectrum of star #22 is characterized by emission in He I 2.059 μ m (weak), Fe II 2.08958 μ m, probable N III 2.11467 μ m, and Br γ (strong). O III / N III is a signature of massive stars from O2 to O8; usually, O4–O7 stars have additional C IV lines, although they are faint in dwarfs. Star #22 appears to still be partially enshrouded; the iron emission, which is indicative of shocks, is located at the position of the star, with diffuse H₂ emission in its surroundings (lines 1–0 (S1), 1–0 (S0), 2–1

Table 2
List of Lines Detected in Early-type Stars

ID	Center (μm)	FWHM (\AA)	EW ^a (\AA)	S/N ^b	V_{LSR}^d (km s^{-1})
M15	2.059024	19 ± 0	-52.1 ± 0.4	278	63 ± 10
M15	2.166333	26 ± 0	-33.6 ± 0.4	355	44 ± 10
1	2.045707	33 ± 0	12.7 ± 1.0	6	...
1	2.111480	229 ± 1	-43.1 ± 5.8	18	...
1	2.163296	199 ± 2	-106.3 ± 16.3	32	...
1	2.188899	195 ± 1	-186.2 ± 14.4	68	...
2	2.165926	57 ± 8	6.3 ± 6.0	7	...
3	2.113018	19 ± 1	1.8 ± 0.9	10	...
3	2.166116	56 ± 0	6.9 ± 1.1	27	...
4	2.112701	38 ± 1	1.8 ± 0.4	12	...
4	2.165608	67 ± 2	3.2 ± 1.3	10	...
5 ^c	2.165962	77 ± 187	2.3 ± 7.1	4	...
6	2.112678	17 ± 11	2.0 ± 1.7	4	...
6	2.165574	73 ± 8	4.1 ± 2.7	6	...
7	2.058728	21 ± 4	0.7 ± 1.1	5	...
7	2.069326	11 ± 4	-0.2 ± 0.5	4	...
7	2.079126	11 ± 1	-0.5 ± 0.5	11	...
7	2.115881	32 ± 0	-1.5 ± 0.8	8	...
7	2.165977	55 ± 2	3.5 ± 1.2	13	...
7	2.189267	16 ± 14	1.9 ± 0.5	11	...
8 ^c	2.069469	13 ± 21	-0.4 ± 0.7	4	...
8	2.079379	16 ± 0	-1.4 ± 0.6	12	...
8	2.115700	46 ± 0	-4.3 ± 0.7	15	...
8	2.189147	15 ± 8	0.6 ± 0.9	4	...
8	2.427893	38 ± 11	-2.1 ± 2.0	4	...
9	2.166162	65 ± 34	3.4 ± 3.9	6	...
10	2.112934	17 ± 1	0.5 ± 1.3	7	...
10	2.166113	92 ± 6	6.9 ± 1.8	14	...
11	2.058704	13 ± 4	0.4 ± 2.1	4	...
11	2.079061	20 ± 3	-1.4 ± 0.8	8	...
11	2.115609	38 ± 3	-2.3 ± 0.7	12	...
11	2.166673	34 ± 6	2.6 ± 1.5	11	...
11	2.189657	20 ± 11	1.2 ± 1.1	6	...
11 ^c	2.427892	17 ± 19	0.0 ± 2.7	2	...
12	2.165692	57 ± 16	8.9 ± 5.1	8	...
13	2.112964	21 ± 15	-0.3 ± 1.1	3	...
13	2.166270	54 ± 3	5.9 ± 1.4	15	...
14	2.166573	...	1.3 ± 3.6	6	...
15	2.166294	81 ± 30	7.6 ± 3.9	8	...
16	2.166464	61 ± 0	10.4 ± 2.4	18	...
17	2.112767	10 ± 7	0.3 ± 1.2	6	...
17	2.165805	78 ± 4	5.8 ± 2.1	11	...
18	2.112785	19 ± 1	0.3 ± 0.4	14	...
18	2.165775	84 ± 0	5.6 ± 1.4	13	...
19	2.166143	44 ± 15	-0.1 ± 4.6	5	...
20	2.166230	92 ± 24	11.6 ± 6.9	6	...
21	2.166324	71 ± 0	8.0 ± 2.2	14	...
22	2.166032	16 ± 0	-5.7 ± 2.5	21	...
23	2.058749	16 ± 0	0.8 ± 2.0	6	...
23 ^c	2.068678	7 ± 420	0.5 ± 1.1	3	...
23	2.079160	16 ± 11	-0.5 ± 0.5	6	...
23	2.115583	25 ± 3	...	10	...
23	2.165405	96 ± 23	3.5 ± 1.4	11	...
23	2.189342	8 ± 6	1.4 ± 0.6	10	...
23	2.427330	27 ± 8	-1.2 ± 1.5	5	...

Notes.

^a Errors on the EWs are calculated following Vollmann & Eversberg (2006).

^b S/N = flux(peak)/continuum noise.

^c Marked ID numbers indicate hints for lines (with a peak S/N $\gtrsim 2$) with poor measurements.

^d The used SINFONI setting allows for an absolute wave calibration within 10 km s^{-1} . For star M15, the average offset of detected OH lines from their rest wavelengths is 3 km s^{-1} , with $\sigma = 8 \text{ km s}^{-1}$. Quoted errors are sqrt (centererr²+10²).

(S1), 1–0 (Q1), 1–0 (Q2), and 1–0 (Q3); Black & van Dishoeck 1987; Gautier et al. 1976; Scoville et al. 1983).

The spectrum of star #14 shows only a Br_γ line in absorption with a central emission peak. Stars #14 and #22 have spectral morphologies that are reminiscent of stars associated with ultracompact H II regions (Bik et al. 2005, 2006) and, pre-empting Section 4.2, IR excesses that are suggestive of emission from natal circumstellar envelopes. We therefore conclude that both are likely to be very young, early-type stars.

The spectra of stars #3, #4, #6, #10, #13, #17, and #18 show both He I 2.112 μm and Br_γ in absorption, indicative of spectral types B0 to B5. The spectra of stars #2, #5, #9, #12, #15, #16, #19, #20, and #21 are noisy; however, a Br_γ line in absorption is clearly visible. They have spectral types earlier than G types.

4.1.2. The Candidate LBV [MFD2011] 15

During the spectroscopic campaign, we re-observed [MFD2011] 15—the LBV candidate #15 in the cluster Cl 1813–178 (Messineo et al. 2011)—in order to search for the spectroscopic variability characteristic of this phase of stellar evolution.

[MFD2011] 15 was first identified with NIRPEC observations (McLean et al. 1998) with $R = 1900$ by Messineo et al. (2011); our new SINFONI spectrum benefits from twice the spectral resolution and an improved signal-to-noise ratio and is shown in Figure 4. Comparison to the earlier spectrum shows an almost identical morphology, with a strong P Cygni line in He I at 2.059 μm , as well as single-peaked emission in He I/ N III/ C III at 2.11407 μm , Mg II at 2.13764 μm and 2.14411 μm , Br_γ at 2.16655 μm , and He I at 2.185 μm . The larger baseline of the SINFONI detector and higher resolving power allow detections of H I at 1.94552 μm , He I at 1.95556 μm , a line emission at 2.10224 μm (likely due to Si III 8–7), and a forest of Pfund H lines; intriguingly, we also detected Si IV at 2.42724 μm . With the exception of the Si IV transition, the spectrum of [MFD2011] 15 bears a strong resemblance to that of the bona fide LBV P Cygni (e.g., Clark et al. 2011). The lack of He II lines results in a degeneracy on the temperature estimate (Najarro et al. 1994, 1997; Martins et al. 2007). Nevertheless, the presence of Si IV in emission may point to a higher temperature than previously assumed from quantitative analysis ($\sim 16 \text{ kK}$; Messineo et al. 2011) and, therefore, higher luminosity; this is of interest since the current estimate ($\log (L/L_\odot) \sim 5.3$) places the star among the faintest of known (candidate) LBVs (Clark et al. 2009a). [MFD2011] 15 is included in a sample of Galactic LBVs that is homogeneously remodeled (M. del Mar Rubio-Diez et al. 2015, in preparation).

4.1.3. Late-type Stars

K-band spectra of late-type stars are characterized by CO bands in absorption with a band-head at 2.2935 μm . We corrected these spectra for interstellar extinction by using the extinction law by Messineo et al. (2005) and the J , H , and K_s magnitudes given in Table 6. We measured the CO equivalent widths from 2.290 to 2.320 μm with a continuum from 2.285 to 2.290 μm , as in Figer et al. (2006). The initial assumption of an average intrinsic $J-K_s = 1.05 \text{ mag}$ ($H-K_s = 0.23 \text{ mag}$) introduces an uncertainty in A_{K_s} of only 0.075 mag, which is

Table 3
Detected *G*-, *K*-, and *M*-type Giants

ID	R.A.[J2000]	Decl.[J2000]	EW(CO)	Sp ^a RSG	Sp ^a RGB	Obs. Date	ID	R.A.[J2000]	Decl.[J2000]	EW(CO)	Sp ^a RSG	Sp ^a RGB	Obs. Date
24	18 14 18.4	-17 56 18	22.4 ± 5.4	<K1	K1	2011 May 29	59	18 13 53.8	-17 57 19	64.9 ± 3.1	M3	...	2011 May 29
25	18 13 59.1	-17 27 31	48.6 ± 1.6	K5	M7	2011 Aug 12	60	18 13 53.8	-17 57 18	21.3 ± 1.9	<K1	K1	2011 May 29
26	18 14 07.8	-17 29 44	32.3 ± 1.7	K1	M0	2011 Aug 19	61	18 13 55.3	-17 57 03	57.2 ± 1.2	M1	...	2011 May 29
27	18 14 02.0	-17 27 24	26.0 ± 1.8	<K1	K3	2011 Aug 20	62	18 13 55.4	-17 57 41	52.9 ± 1.9	M0	...	2011 May 29
28	18 10 52.0	-17 42 23	44.5 ± 1.1	K4	M5	2011 Jun 24	63	18 13 52.3	-17 56 51	48.4 ± 1.4	K5	M7	2011 May 29
29	18 10 51.8	-17 42 20	25.9 ± 19.8	<K1	K4	2011 Jun 24	64	18 13 52.2	-17 57 56	36.5 ± 3.2	K2	M2	2011 May 29
30	18 10 58.3	-17 41 24	53.2 ± 0.8	M0	...	2011 Aug 19	65	18 13 52.3	-17 57 54	49.5 ± 1.0	K5	M7	2011 May 29
31	18 13 54.9	-17 53 49	45.4 ± 6.8	K5	M6	2011 Jun 29	66	18 13 52.0	-17 57 52	29.0 ± 8.1	K1	K5	2011 May 29
32	18 13 44.9	-17 57 12	33.8 ± 2.8	K2	M1	2011 Aug 19	67	18 14 00.1	-17 56 31	41.4 ± 2.0	K3	M4	2011 May 29
33	18 13 50.2	-17 54 26	50.3 ± 1.2	M0	M7	2011 May 31	68	18 14 00.8	-17 56 51	43.2 ± 1.0	K4	M5	2011 May 29
34	18 14 03.1	-17 58 52	18.8 ± 2.2	<K1	<K1	2011 Aug 20	69	18 14 01.0	-17 56 50	21.6 ± 3.9	<K1	K1	2011 May 29
35	18 14 03.2	-17 58 50	44.0 ± 3.8	K4	M6	2011 Aug 20	70	18 14 09.4	-18 00 32	37.6 ± 0.6	K2	M2	2011 May 29
36	18 14 03.4	-17 58 49	42.0 ± 3.5	K4	M5	2011 Aug 20	71	18 10 52.5	-17 41 11	51.2 ± 0.8	M0	M7	2011 Aug 19
37	18 14 02.7	-17 55 38	49.0 ± 2.2	K5	M7	2011 Jun 24	72	18 13 52.7	-17 58 03	39.7 ± 2.9	K3	M3	2011 Aug 12
38	18 13 49.9	-17 57 05	43.7 ± 2.0	K4	M5	2011 Aug 19	73	18 13 52.9	-17 58 02	27.1 ± 2.4	<K1	K3	2011 Aug 12
39	18 14 27.7	-17 57 05	30.9 ± 1.6	K1	K5	2011 Jun 24	74	18 14 08.5	-18 00 19	20.9 ± 7.3	<K1	K1	2012 Jun 26
40	18 14 06.3	-17 28 33	18.9 ± 1.7	<K1	<K1	2011 Jun 24	75	18 14 08.8	-18 00 17	21.2 ± 3.5	<K1	K1	2012 Jun 26
41	18 14 05.6	-17 28 50	49.8 ± 1.1	K5	M7	2011 Jun 24	76	18 14 09.4	-18 00 48	32.4 ± 1.8	K1	M0	2012 Aug 13
42	18 10 54.8	-17 39 56	40.9 ± 2.2	K3	M4	2011 Jun 07	77	18 13 49.1	-17 56 15	25.0 ± 1.3	<K1	K2	2012 Sep 15
43	18 10 55.1	-17 40 25	31.8 ± 2.7	K1	K5	2011 Jun 07	78	18 13 55.4	-17 54 31	24.8 ± 1.6	<K1	K2	2012 Sep 02
44	18 10 52.7	-17 40 19	47.1 ± 1.1	K5	M7	2011 Jun 07	79	18 13 53.4	-17 55 19	21.2 ± 2.6	<K1	K1	2012 Sep 02
45	18 10 52.7	-17 40 08	44.1 ± 1.4	K4	M5	2011 Jun 07	80	18 13 52.2	-17 56 22	42.4 ± 0.9	K4	M4	2012 Aug 08
46	18 10 50.5	-17 40 29	23.7 ± 2.4	<K1	K2	2011 Jun 07	81	18 13 57.2	-17 58 08	31.9 ± 2.6	K1	M0	2012 Aug 11
47	18 10 55.2	-17 41 20	62.5 ± 9.0	M2	...	2011 May 20	82	18 13 57.2	-17 58 06	22.8 ± 6.6	<K1	K2	2012 Aug 11
48	18 10 56.6	-17 41 54	44.4 ± 1.6	K4	M5	2011 Jun 07	83	18 13 38.1	-17 43 19	46.5 ± 0.8	K5	M6	2012 Jun 03
49	18 14 01.8	-17 54 43	46.7 ± 1.2	K5	M7	2011 Aug 12	84	18 13 20.8	-18 06 26	44.6 ± 0.6	K4	M5	2012 Jun 21
50	18 14 03.9	-17 55 06	52.2 ± 1.5	M0	...	2011 Aug 12	85	18 13 13.5	-17 48 07	54.6 ± 0.7	M0	...	2012 Jun 03
51	18 13 54.7	-17 54 57	56.1 ± 1.4	M1	...	2011 Jun 24	86	18 14 44.5	-18 07 38	50.2 ± 1.3	K5	M7	2012 Jun 21
52	18 13 53.4	-17 55 10	14.2 ± 2.5	<K1	<K1	2011 Jun 24	87	18 13 48.2	-17 50 42	57.5 ± 1.8	M1	...	2012 Jun 21
53	18 13 52.9	-17 55 00	53.4 ± 1.4	M0	...	2011 Jun 24	88	18 13 54.8	-18 06 56	49.9 ± 2.6	M0	M7	2012 Jun 21
54	18 13 52.5	-17 56 16	60.0 ± 1.8	M2	...	2011 May 29	89	18 14 07.8	-17 28 37	55.4 ± 2.8	M1	...	2012 Sep 15
55	18 13 51.9	-17 56 27	44.9 ± 1.1	K4	M6	2011 May 29	90	18 14 02.9	-17 29 01	40.0 ± 1.2	K3	M4	2012 Sep 15
56	18 13 55.5	-17 56 18	43.4 ± 0.4	K4	M5	2011 May 29	91	18 14 05.5	-17 29 25	30.9 ± 3.0	K1	K5	2012 Aug 19
57	18 13 54.6	-17 56 12	57.1 ± 7.0	M1	...	2011 May 29	92	18 14 10.1	-17 27 57	29.4 ± 2.1	K1	K4	2012 Sep 15
58	18 13 54.1	-17 57 22	40.8 ± 6.5	K3	M4	2011 May 29	93	18 13 47.4	-17 57 10	41.3 ± 1.6	K3	M4	2012 Jun 27

Note.

^a Spectral types are estimated by using the relation between spectral types and EW(CO)s of red giants (RGBs), as well as that between spectral types and EW(CO)s of RSGs.

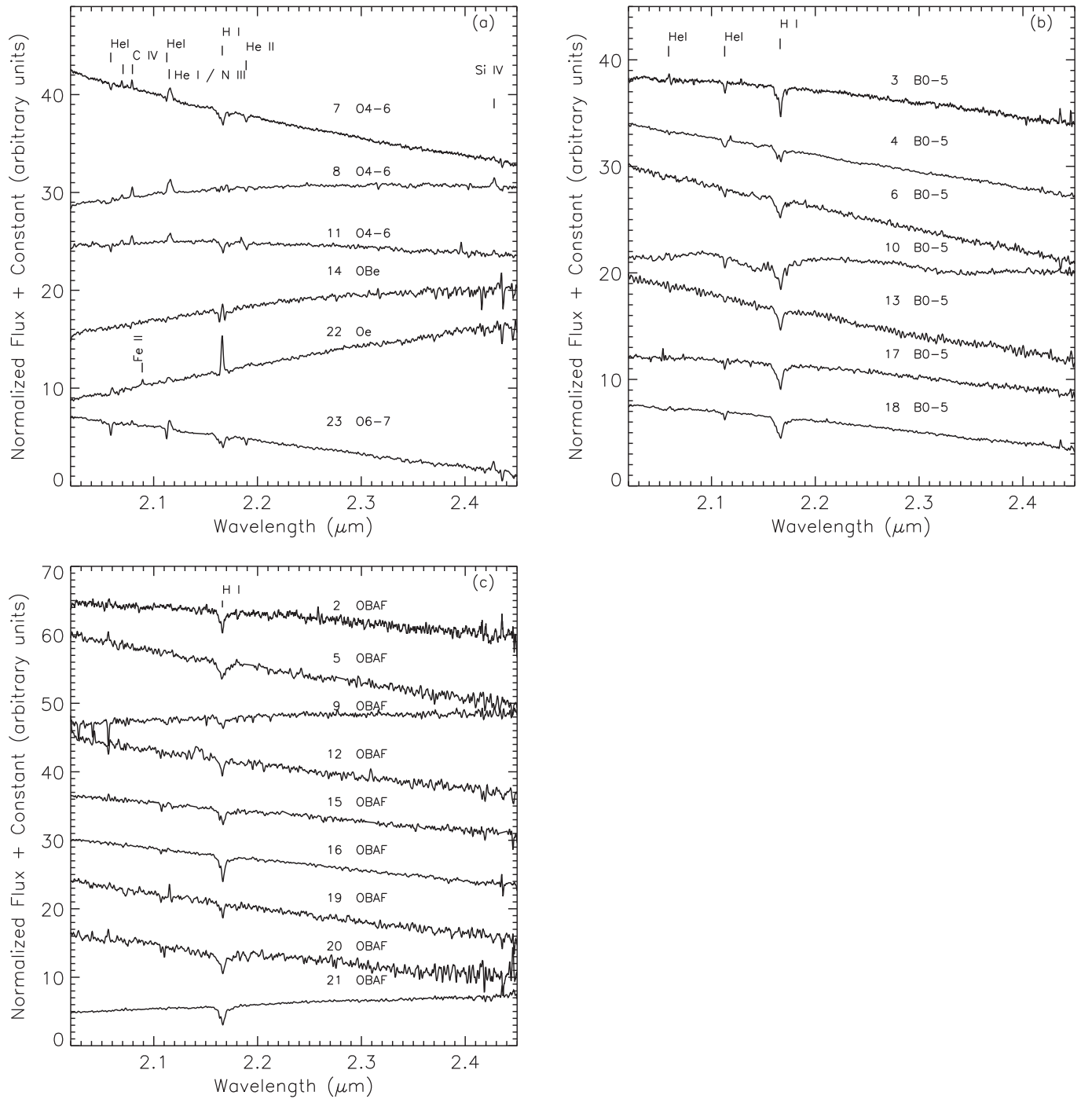


Figure 3. SINFONI K -band spectra of newly detected O-type stars. Panel (a) displays massive O-type stars, panel (b) B-type stars, and panel (c) spectra (mostly marginal detections) with a detection of a Br_γ line.

negligible for spectral classification. A shift by 10% on the K -band extinction produces a median shift in EW of 1.4%. We estimated the uncertainty due to the continuum region adopted by adding small shifts and remeasuring; the percentage uncertainty has a median value of 5%. The list of detected late types is provided in Table 3.

Spectral types were obtained by comparison with template spectra of red giants and red supergiants (Kleinmann & Hall 1986). For each star, spectral types are provided for two possible luminosity classes—giants and supergiants—which

follow differing relations between $\text{EW}(\text{CO})$ s and spectral types (e.g., Figer et al. 2006; Messineo et al. 2014a). Typically, the spectral types obtained via this methodology are accurate to within two spectral types. For 13 stars in Table 3 we found $\text{EW}(\text{CO})$ s larger than 52 \AA , i.e., larger than those of an M7 III; the spectra of seven of them (#30, #47, #50, #54, #57, #87, and #89) show water absorption and are likely Mira asymptotic giant branch (AGB) stars (see, e.g., Messineo et al. 2014b); the spectra of stars #51, #53, and #85 are quite noisy; the spectra of stars #59, #61, and #62 do not show water

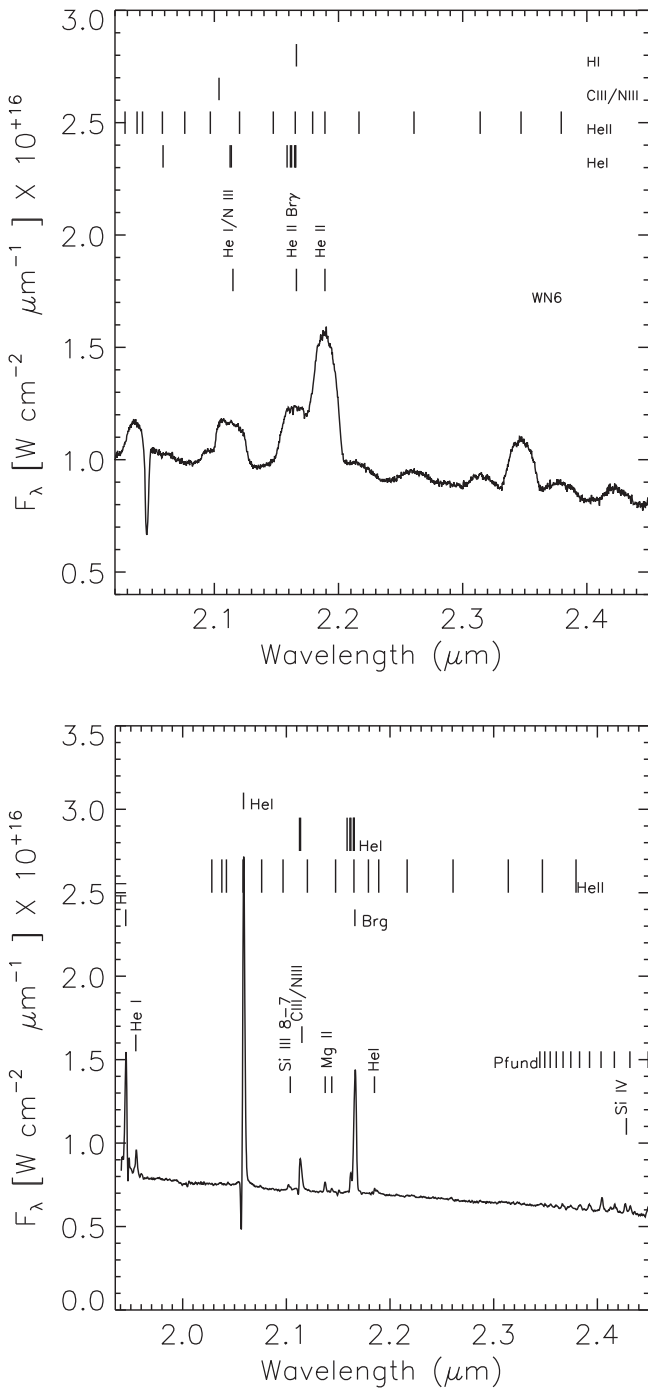


Figure 4. Upper panel: SINFONI K -band spectrum of a newly detected WN6 star. Lower panel: SINFONI K -band spectrum of the candidate LBV [MFD2011] 15.

absorption and could be semi-regular AGB stars (Messineo et al. 2014b).

4.2. Extinction in K_s Band and Bolometric Corrections

For early-type stars, we assumed intrinsic colors and effective temperatures, T_{eff} , as tabulated per spectral type in Messineo et al. (2011), and based on the works by Bibby et al. (2008), Crowther et al. (2006b), Johnson (1966), Koornneef (1983), Humphreys & McElroy (1984), Lejeune & Schaerer

(2001), Martins et al. (2005), Martins & Plez (2006), and Wegner (1994). For the WN6, we used the T_{eff} values and average infrared colors listed by Crowther (2007) and Crowther et al. (2006a). For late-type stars, we adopted the intrinsic colors given by Koornneef (1983).

Total extinction in K_s band was calculated by assuming these intrinsic colors and by adopting the power-law curve with an index of -1.9 by Messineo et al. (2005). Estimates for the IR excess in three different colors ($E(J-H)$, $E(H-K_s)$, and $E(J-K_s)$) are provided in Table 4. Since the K_s band of mass-losing early-type stars (e.g., WR) may have significant excess due to free-free emission and dust (Cohen et al. 1975), it is preferable to use $E(J-H)$. For late-type stars, the $E(J-K_s)$ is typically used.

Apparent bolometric magnitudes are calculated with dereddened K_s magnitudes and bolometric corrections, BC_K , as listed in Tables 8, 9, and 10 by Messineo et al. (2011, and references therein); for the WN6, the adopted BC_K is taken from Crowther et al. (2006a); for late-type stars BC_K values per spectral type are available from the work of Levesque et al. (2005). As shown in the $J-H$ versus $H-K_s$ diagram and in the K_s-8 versus $H-K_s$ diagram of Figure 5, the bulk of the sources follow the direction expected for reddening by interstellar dust. The two early-type stars #14 (OBe) and #22 (Oe) show significant infrared excess, possibly due to the presence of circumstellar material.

5. STELLAR PARAMETERS OF W33 MEMBERS

In the following, we discuss the properties of the detected massive early-type stars and their association with W33. Recently, Immer et al. (2013) measured parallactic distances of several water masers in the direction of W33. The centroid LSR velocities of three out of four H_2O maser sites are between 29 and 37 km s^{-1} , while that of the remaining, W33B, is 59.3 km s^{-1} . Nevertheless, the trigonometric parallaxes yield similar values for their distances.

Their average distance is 2.64 kpc with a standard deviation $\sigma = 0.25$ kpc. Following Immer et al., for the entire W33 complex we adopt the parallactic distance to the maser W33B of $2.4^{+0.17}_{-0.15}$ kpc ($DM = 11.90 \pm 0.16$ mag). Derived absolute K_s , M_K , and bolometric magnitudes M_{bol} are listed in Tables 4 and 5.

5.1. O-type Stars

We detected three luminous O stars within W33 as part of our survey—#7, #8 (O4–6), and #23 (O6–7)—which are among the brightest stars appearing in the K_s versus $H-K_s$ diagrams of Figure 8.

Stars #7 and #8 are located in the Mercer1 region, which coincides with the H II region G12.745–00.153 (Lockman 1989; White et al. 2005)—in SIMBAD this object is named [L89b] 12.745–00.153. The two O-type stars have an angular separation of $1/3$; the first has an $A_{K_s} = 0.82 \pm 0.05$ mag, the latter has an A_{K_s} more than 3 times larger. The location of star #7 near the peak of the 24 μm emission of the H II region (Carey et al. 2009) provides evidence for its association with W33; star #8 is located in the dustier surrounding 8 μm shell (visible in GLIMPSE). Similar variations of A_{K_s} are reported within other mid-IR bubbles (e.g., Bik et al. 2010); therefore, we attribute the difference of interstellar extinction between the two stars to strong dust variations of the same H II in W33. Star

Table 4
Color Properties of Newly Detected Early-type Stars

ID	Sp. Type	$(J - K_s)_o$ (mag)	$(H - K_s)_o$ (mag)	$A_{K_s}(JH)$ (mag)	$A_{K_s}(JK_s)$ (mag)	$A_{K_s}(HK_s)$ (mag)	$Q1$ (mag)	$Q2$ (mag)
1	WN6	0.370	0.260	0.981	0.953	0.902	-0.272 ± 0.075	-2.784 ± 0.008
2	OBAF	-0.050	-0.040	1.123	1.093	1.041	0.149 ± 0.043	...
3	B0-5	-0.160	-0.080	1.470	1.353	1.144	0.442 ± 0.137	...
4	B0-5	-0.160	-0.080	1.267	1.215	1.123	0.224 ± 0.077	-1.017 ± 0.086
5	OBAF	-0.050	-0.040	0.337	0.375	0.444	-0.071 ± 0.045	...
6	B0-5	-0.160	-0.080	1.254	1.295	1.369	-0.086 ± 0.658	...
7	O4-6	-0.210	-0.100	0.819	0.803	0.775	0.116 ± 0.163	0.318 ± 0.010
8	O4-6	-0.210	-0.100	2.883	2.809	2.676	0.292 ± 0.138	0.270 ± 0.143
9	OBAF	-0.050	-0.040	2.722	2.682	2.612	0.168 ± 0.043	...
10	B0-5	-0.160	-0.080	1.978	1.972	1.961	0.065 ± 0.051	...
11	O4-6	-0.210	-0.100	2.280	2.240	2.169	0.182 ± 0.107	...
12	OBAF	-0.050	-0.040	0.313	0.243	0.117	0.296 ± 0.054	...
13	B0-5	-0.130	-0.030	1.067	1.029	0.962	0.071 ± 0.059	...
14	OBe	-0.120	-0.060	2.465	2.622	2.902	-0.502 ± 0.061	...
15	OBAF	-0.060	-0.010	0.658	0.560	0.386	0.289 ± 0.054	...
16	OBAF	-0.060	-0.010	0.580	0.518	0.408	0.169 ± 0.063	...
17	B0-5	-0.130	-0.030	1.832	1.831	1.828	-0.059 ± 0.045	...
18	B0-5	-0.130	-0.030	1.657	1.639	1.608	-0.003 ± 0.051	...
19	OBAF	-0.050	-0.040	0.607	0.511	0.340	0.378 ± 0.051	...
20	OBAF	-0.050	-0.040	0.529	0.491	0.423	0.184 ± 0.048	...
21	OBAF	-0.050	-0.040	0.602	0.591	0.570	0.095 ± 0.046	...
22	Oe	-0.210	-0.100	2.872	3.085	3.465	-0.670 ± 0.129	-6.532 ± 0.344
23	O6-7	-0.210	-0.100	1.199	1.165	1.104	0.173 ± 0.083	...

Notes. The $Q1$ and $Q2$ parameters are defined as in Messineo et al. (2012).

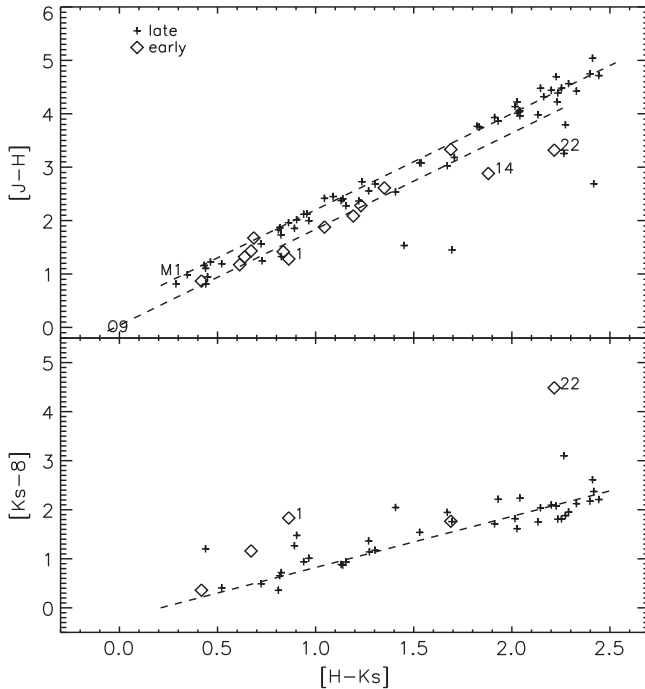


Figure 5. Top panel: $H-K_s$ vs. $J-H$ colors of observed stars are shown. The two dashed lines trace the locus of an M1 star (upper line) and O9 star (lower line) with increasing extinction A_{K_s} from 0.0 mag to 3.5 mag. Bottom panel: $H-K_s$ vs. K_s-8 colors. The dashed line traces the locus of a star ($H-K_s = 0$ mag, $K_s-8 = 0$ mag) with increasing extinction A_{K_s} from 0.0 mag to 3.5 mag.

#23 is located in the c11 region and has an A_{K_s} value of 1.20 ± 0.03 mag (similar to that of star #7).

By assuming a common distance of 2.4 kpc, bolometric corrections as listed in Table 5, and a solar bolometric constant

of -4.74 mag, we derived $\log(L/L_\odot) = 5.51^{+0.09}_{-0.09}$, $5.61^{+0.09}_{-0.09}$, and $5.56^{+0.07}_{-0.07}$ and $M_K = -4.63 \pm 0.17$ mag, -4.89 ± 0.18 mag, and -5.00 ± 0.16 mag for stars #7, #8, and #23, respectively. The similarity of M_{bol} (and M_K) values suggests that these three stars have similar ages and masses. Supporting our assertions in the preceding section regarding their likely evolved nature, comparison with M_K values of mid- O-type stars (Martins & Plez 2006) indicates that all three are consistent with luminosity classes III–I.

By using the latest stellar models by the Geneva group with solar metallicity and rotation (Ekström et al. 2012), we derive stellar masses from 30 to 40 M_\odot and an age below 6 Myr, at which point all 40 M_\odot stars would have been lost to supernovae (SNe); the presence of spectroscopically O4–6 supergiants ($M_K = -5.03$ mag with $\sigma = 0.47$ mag) within the Arches (2–4 Myr; Martins et al. 2008) and of O4–6 dwarfs ($M_K = -4.15$ mag with $\sigma = 0.43$ mag) in Danks 1 ($\sim 1.5^{+1.5}_{-0.5}$ Myr; Davies et al. 2012) suggests an age of 2–4 Myr.

At the position of Oe star #22, the SINFONI cube shows several H_2 lines, suggesting that star #22 ($K_s = 10.448$ mag) is still embedded. We derived $A_{K_s} = 2.87 \pm 0.07$ mag from the $J-H$ color excess; by assuming a distance of 2.4 kpc, we estimate $M_K = -4.33 \pm 0.18$ mag. For confirming its luminosity class, further spectroscopy in J and H band is required.

5.2. A New WN6

Found in the southwest periphery of W33 (as shown in Figure 7), the WN6b star #1 has an $A_{K_s} = 1.0$ mag, when

¹⁵ The average M_K is calculated with the magnitudes from Figer et al. (2002), a distance of 8.4 kpc, and the extinction law by Messineo et al. (2005).

¹⁶ The average M_K is calculated with the magnitudes and distance provided by Davies et al. (2012) and the extinction law by Messineo et al. (2005).

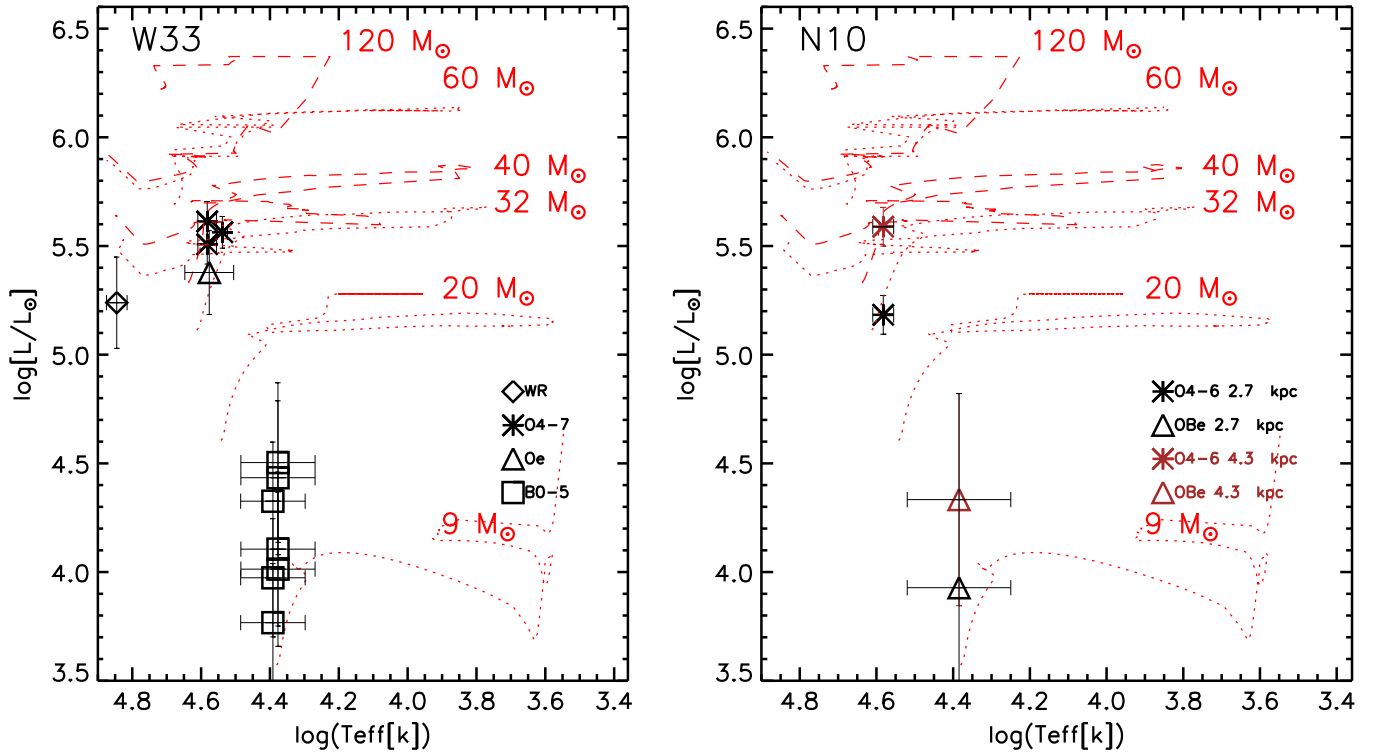


Figure 6. Luminosity vs. T_{eff} diagram of massive O-type stars (O4–6 and O6–7), B0–5 stars, and the WN6b star detected in the direction of W33 (left) and of N10 (right). For W33, we assume a distance of 2.4 kpc; for N10, the same two stars are plotted for a distance of 2.4 kpc (black points) and for a distance of 4.3 kpc (lighter points). Stellar tracks for stars of 9, 20, 32, 40, 60, and 120 M_{\odot} , with solar metallicity and including rotation, are shown with dotted and dashed curves (Ekström et al. 2012).

Table 5
Physical Properties of Newly Detected Massive Stars

ID ^a	Sp.	Class ^b	T_{eff} (K)	K_{S_0} (mag)	A_{K_S} (mag)	BC_{K_S} (mag)	$M_{K_S}(\text{der.})$ (mag)	$\log(L/L_{\odot})$ (mag)	Region
1	WN6	I	70000 ± 5000	7.043 ± 0.033	0.981 ± 0.025	-3.50 ± 0.50	-4.86 ± 0.16	$5.24^{+0.21}_{-0.21}$	South of W33
3	B0–5	III	24000 ± 6700	9.437 ± 0.059	1.470 ± 0.040	-2.83 ± 0.87	-2.46 ± 0.17	$4.01^{+0.35}_{-0.35}$	Mercer1, W33
4	B0–5	III	24000 ± 6700	8.386 ± 0.035	1.267 ± 0.025	-2.83 ± 0.87	-3.52 ± 0.16	$4.43^{+0.35}_{-0.35}$	Mercer1, W33
6	B0–5	III	24000 ± 6700	8.212 ± 0.245	1.254 ± 0.208	-2.83 ± 0.87	-3.69 ± 0.29	$4.50^{+0.37}_{-0.37}$	Mercer1, W33
7	O4–6	I	38000 ± 2500	7.271 ± 0.057	0.819 ± 0.051	-4.40 ± 0.15	-4.63 ± 0.17	$5.51^{+0.09}_{-0.09}$	Mercer1, W33
8	O4–6	I	38000 ± 2500	7.009 ± 0.067	2.883 ± 0.058	-4.40 ± 0.15	-4.89 ± 0.17	$5.61^{+0.09}_{-0.09}$	Mercer1, W33
10	B0–5	III	24000 ± 6700	9.207 ± 0.020	1.978 ± 0.015	-2.83 ± 0.87	-2.69 ± 0.16	$4.11^{+0.35}_{-0.35}$	Mercer1, W33
13	B0–5	V	25000 ± 5900	9.827 ± 0.026	1.067 ± 0.016	-3.12 ± 0.66	-2.07 ± 0.16	$3.97^{+0.27}_{-0.27}$	Mercer1, W33
17	B0–5	V	25000 ± 5900	10.343 ± 0.021	1.832 ± 0.013	-3.12 ± 0.66	-1.56 ± 0.16	$3.77^{+0.27}_{-0.27}$	cl2, W33
18	B0–5	V	25000 ± 5900	8.946 ± 0.023	1.657 ± 0.015	-3.12 ± 0.66	-2.96 ± 0.16	$4.33^{+0.27}_{-0.27}$	cl2, W33
22	Oe	7.576 ± 0.075	2.872 ± 0.071	...	-4.33 ± 0.18	...	cl1, W33
23	O6–7	I	35000 ± 1200	6.904 ± 0.037	1.199 ± 0.027	-4.17 ± 0.08	-5.00 ± 0.16	$5.56^{+0.07}_{-0.07}$	cl1, W33
11	O4–6	I	38000 ± 2500	8.083 ± 0.039	2.280 ± 0.032	-4.40 ± 0.15	-5.08 ± 0.10	$5.69^{+0.07}_{-0.07}$	BDS2003–115
14	OBe	9.691 ± 0.028	2.465 ± 0.026	...	-3.47 ± 0.10	...	BDS2003–115

Notes. A distance of $2.4^{+0.17}_{-0.15}$ kpc ($DM = +11.90 \pm 0.16$ mag) is used for W33 (Immer et al. 2013) and of $4.29^{+0.17}_{-0.19}$ kpc ($DM = 13.16 \pm 0.09$ mag) for N10.

^a OBAF detections are not included in this table.

^b Classes are photometrically estimated using M_K values from Martins & Plez (2006), Bibby et al. (2008), Humphreys & McElroy (1984), and Lejeune & Schaerer (2001).

assuming the average intrinsic near-infrared colors for late WN stars by Crowther et al. (2006a). With a distance of 2.4 kpc we measured $M_K = -4.86 \pm 0.16$ mag, which fits well with the average $M_K = -5.13$ mag ($\sigma = 0.07$ mag) of two other WN6b stars analyzed by Crowther et al. (2006a). By using $BC_K = -3.5 \pm 0.5$ mag, we obtain $\log(L/L_{\odot}) = 5.24^{+0.21}_{-0.21}$ and a mass

of about $27 \pm 2.5 M_{\odot}$. The detection of a late WN implies the existence of highly luminous progenitor supergiants (Georgy et al. 2012), such as those we detected in Mercer1.

The compilations of Galactic WR stars by van der Hucht (2001), Mauerhan et al. (2011), Shara et al. (2012), and Faherty et al. (2014), identify 43 WN6 stars from a total of 443

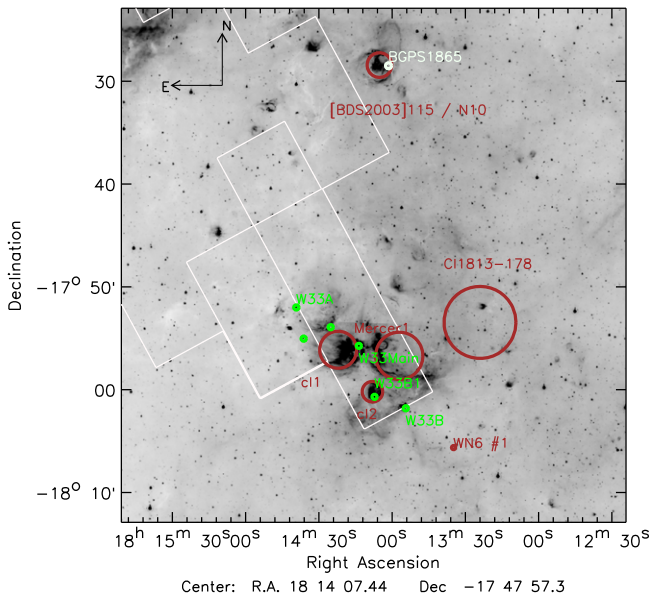


Figure 7. $8\ \mu\text{m}$ emission from GLIMPSE (Churchwell et al. 2009) of W33. The Mercer1, cl1, cl2, and Cl1813–178 regions are marked with circles, as well as the isolated WN6 star #1. The green dots indicate the location of the molecular clumps in W33 detected by Immer et al. (2014), and in blue that in N10 studied by Ma et al. (2013). The white contours show the location of the peak of ^{12}CO emission in the range from 30 to $60\ \text{km s}^{-1}$ from the Galactic survey of Dame et al. (2001). The image has a linear size of $34.5\ \text{pc} \times 35.09\ \text{pc}$ for a distance of 2.4 kpc.

WRs of all flavors. Among the 12 WN6 stars listed in the latter three works, 40% have broad features, suggesting that only 4% of known WRs have a similar classification. Rapid rotation and the presence of a magnetic field have been suggested to explain the broadening of their spectral lines and the flattening of the line peaks (Shenar et al. 2014).

5.3. Spectral Types B0–5

Six B0–5 stars were detected. Stars #3, #4, #6, #10, and #13 are located in the Mercer1 region, with A_{K_s} values of 1.47 ± 0.04 , 1.27 ± 0.03 , 1.25 ± 0.21 , 1.98 ± 0.02 , and 1.07 ± 0.02 mag, respectively. For a distance of 2.4 kpc, their M_K values range from -2.07 to -3.69 mag and suggest a mix of dwarfs and giants (Humphreys & McElroy 1984; Wegner 1994; Lejeune & Schaerer 2001; Martins & Plez 2006) with initial masses from 9 to $15\ M_{\odot}$ (Ekström et al. 2012).

The B0–5 stars #18 and #17 are located in the cl2 cluster. Star #18 is the brightest star in the small cluster. For a distance of 2.4 kpc, it has $A_{K_s} = 1.66 \pm 0.02$ mag and $M_K = -2.96 \pm 0.16$ mag, which suggests a luminosity class V/III and an initial mass of $12 \pm 3\ M_{\odot}$ (Ekström et al. 2012). Star #17, with $A_{K_s} = 1.83 \pm 0.01$ mag and $M_K = -1.56 \pm 0.16$ mag, has a likely initial mass of $10 \pm 2\ M_{\odot}$.

We note that the nine stars generically classified as “early” (i.e., those assigned spectral type OBAF in Table 1) exhibit such an uncertainty in temperatures and hence intrinsic colors that we cannot derive meaningful physical properties for them at this time.

5.4. Late-type Stars

For a distance of 2.4 kpc, the detected late-type stars from Table 3 remain fainter than $M_{\text{bol}} = -5.26$ mag ($\log(L/L_{\odot})$

$= 4.0$). They all have magnitudes consistent with those of giant stars.

6. THE GLOBAL STRUCTURE AND STAR FORMATION HISTORY OF W33

Complementing Figure 1, in which we show the location of (candidate) massive stars in the putative individual clusters, Figure 7 delineates the nominal locations of these regions on a map of the $8\ \mu\text{m}$ emission. Massive stars are found throughout W33, with the richest region being the Mercer 1¹⁷ aggregate to the west of W33 Main. The presence of two O4–6 (super-)giants #7 and #8 suggests a burst of star formation occurred $\sim 2\text{--}4$ Myr ago. A further five early to mid-B-type stars (#3, #4, #6, #10, and #13) have A_{K_s} consistent with those of the O-type stars and bolometric magnitudes typical of dwarfs and/or giants. The remaining three stars (#2, #5, and #9) are of generic early (OBAF) spectral-type.

Immediately to the west of Mercer 1 and the embedded protocluster forming within the radio source W33 Main, we find the O4–6 (super-)giant #23, which demonstrates similar physical properties to stars #7 and #8. In GLIMPSE images, star #23 is surrounded by a yellow curved filament, similar to the mid-infrared bow shocks identified by Povich et al. (2008). The likely young massive Oe star #22 is located in front of the apex of the bowshock between #23 and the radio source W33 Main, consistent with its elevated extinction ($A_{K_s} = 2.87 \pm 0.07$ mag). The lack of further massive stars in this region leads us to conclude that the cl1 region does not delineate a bona fide cluster.

The massive protostar W33A (Davies et al. 2010) is located in the northeast of these regions, while the stellar aggregate cl2 is located to the south. A sequence of reddened stars is detected within the compact nebula of cl2 (Figures 1 and 8); these are found at $J-K_s \approx 1.6$ mag in the color–magnitude diagram shown in Figure 8. The two brightest stars, #17 and #18, are spectroscopic B0–5 types, while the remaining three, #19, #20, and #21, are classified as spectral type OBAF. Radio continuum emission is found in the direction of the cl2 cluster, centered on star #18; we estimate a flux density of 0.57 Jy at 20 cm and 0.64 Jy at 90 cm by using MAGPIS data and an aperture of $35''$. Under the assumption of optically thin thermal emission with an electron temperature, $T_e = 10,000\ \text{K}$, this implies a Lyman continuum photon flux, N_{lyc} , of $10^{47.2}\ \text{s}^{-1}$ (e.g., Martín-Hernández et al. 2003; Rubin 1968; Storey & Hummer 1995). For comparison, an O9.5 V emits a number of N_{lyc} of $10^{47.9}\ \text{s}^{-1}$ and an O9.5 III of $10^{48.4}\ \text{s}^{-1}$ (from the more recent work by Martins et al. 2005). By comparing the results from Martins et al. (2005), Vacca et al. (1996), and Panagia (1973), after having corrected for relative average shifts, we estimate $N_{\text{lyc}} = 10^{47.2}\ \text{s}^{-1}$, $10^{48.0}\ \text{s}^{-1}$, $10^{44.4}\ \text{s}^{-1}$, and $10^{45.1}\ \text{s}^{-1}$ for a B0 V, a B0 III, a B2 V, and a B2 III star, respectively. We find typical uncertainties of 0.2 dex for every N_{lyc} value. Therefore, stars #17 and #18 may already account for the requisite ionizing flux. While early B-type stars with initial masses of 9– $12\ M_{\odot}$ are present in stellar populations with ages ranging up to 30 Myr (Ekström et al. 2012), the nebular emission associated with cl2 suggests a much younger age, likely of only a few Myr.

¹⁷ Synonymous with the H II G12.745–00.153.

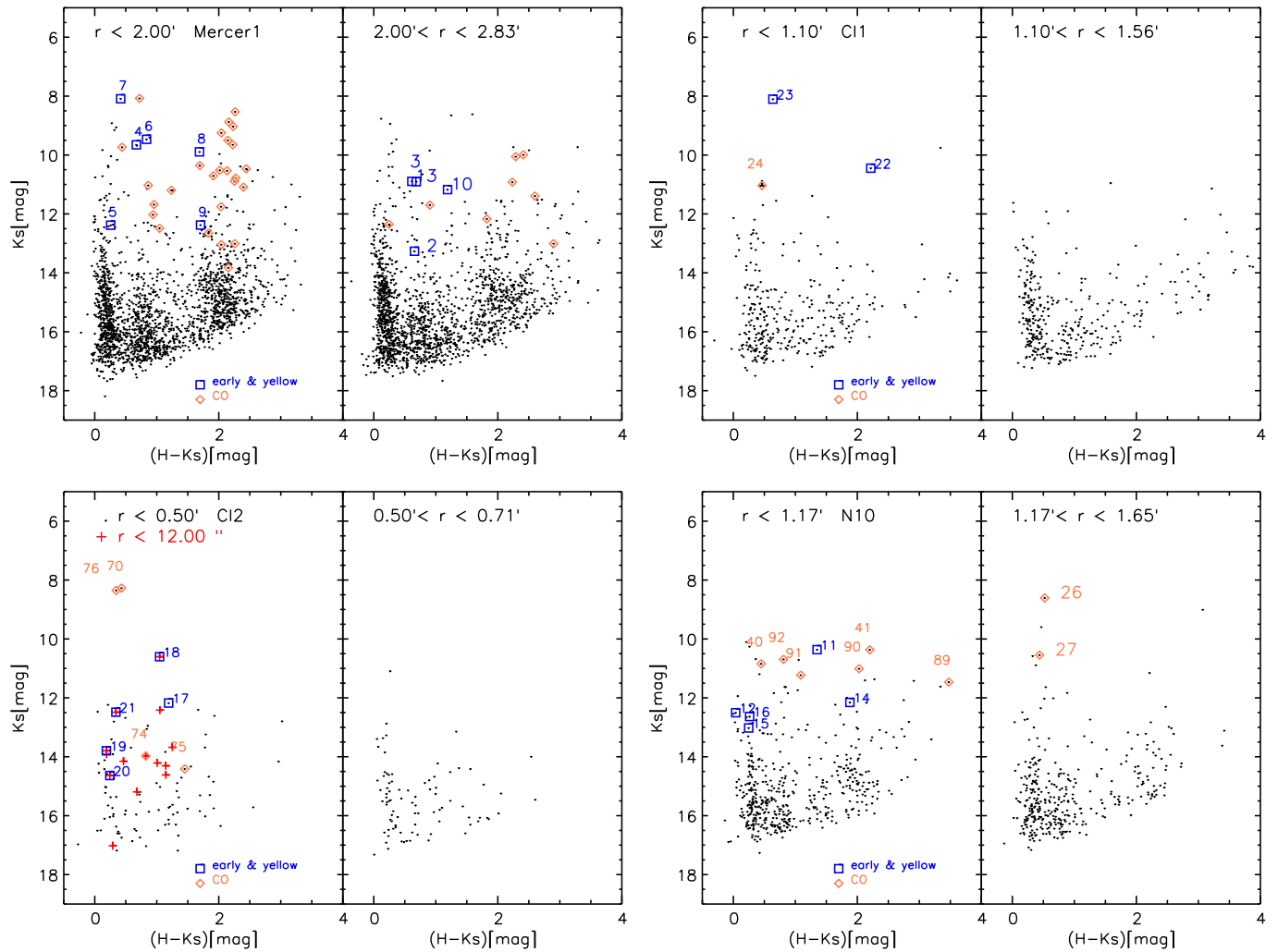


Figure 8. 2MASS-UKIDSS $H-K_s$ vs. K_s diagram of the Mercer1 (top left), cl1 (top right), cl2 (bottom left), and N10 (bottom right) regions. Spectroscopic detected early-yellow stars are indicated with squares, late-type stars with diamonds. Identification numbers are taken from Tables 1 and 3. A comparison field of equal area is shown in the right panel; field datapoints are taken from an annular region. In the cl2 region, stars concentrated in the central 12'' are marked with crosses.

Finally, further to the southwest we find the broad-lined WN6 star #1. We unsuccessfully searched the W33 area for other possible bright stars ($K_s < 10$ mag) with properties of free-free emitters or candidate red supergiants by using the infrared photometric criteria of Messineo et al. (2012). Thus, it is unlikely that there are any further young massive stellar aggregates associated with the complex. Consideration of these findings emphasizes that the massive stellar population is distributed across the confines of the W33 and appears not to be concentrated in rich young clusters similar to, e.g., Danks 1 and Danks 2 within the G305 star-forming complex (Davies et al. 2012).

This behavior appears to mirror the current location of cold molecular material, within which future generations of stars may form. Immer et al. (2014) report the detection of six molecular clumps along the east side of W33, with masses of $(0.2-4.0) \times 10^3 M_\odot$ coincident with the peak of the CO intensity map (see Figure 7). By contrast, the evolved H II region G12.745-00.153 (Mercer1) resides on the west side of W33, where the molecular matter has already been swept out, a configuration that is at least suggestive of sequential star formation. Similarly, the dense clump W33B1 (Immer

et al. 2014) is located $\sim 35''$ southwest of the cl2 cluster on the periphery of the apparent wind-blown nebula associated with the latter; it is conceivable that its influence is contributing to this new protostellar core by heating and compressing it (Immer et al. 2014).

7. BDS2003-115 AND BUBBLE N10

The cluster candidate BDS2003-115 is located north of W33 (Bica et al. 2003; Messineo et al. 2011). It appears as a group of bright near-infrared stars in the core of a mid-infrared bubble (bubble N10, and candidate SN remnant (SNR) G13.1875+0.0389, Churchwell et al. 2006; Helfand et al. 2006; Watson et al. 2008)¹⁸. Radio line observations of

¹⁸ The center of the bubble is dominated by radio continuum emission and $8 \mu\text{m}$ and $24 \mu\text{m}$ emission, the latter likely derived from warm dust not yet destroyed by stellar feedback. With MAGPIS data (Helfand et al. 2006), we estimated an integrated flux density of 5.3 Jy at 20 cm and of 7.5 Jy at 90 cm, over identical areas of $2'$ radii. The resulting spectral index $\alpha \sim -0.2$. Such a value is marginally consistent with that expected from the optically thin free-free emission ($\alpha \sim -0.1$) expected for an H II region, although one cannot exclude an additional source of non-thermal emission, originating from either the shocked stellar winds from massive OB stars or an SN explosion (Williams & Sun 1996; Leitherer et al. 1997; Sidarin et al. 2014).

Table 6
Average M_K , A_{K_s} , and Number of Stars per Spectral Group and Luminosity Class in Cl 1813–178 (Table 2, Messineo et al. 2011)

Sp. Group	Class ^c	Cluster Distance ^a 4.8 kpc			W33 Distance ^b 2.4 kpc		
		M_K^d	$A_{K_s}^d$	Nstar	M_K^d	$A_{K_s}^d$	Nstar
B0–B3	I	-6.32 ± 1.01	0.94 ± 0.25	7	-5.78 ± 0.33	0.87 ± 0.18	3
B0–B3	V/III	-4.24 ± 0.60	0.98 ± 0.34	5	-3.32 ± 0.90	0.97 ± 0.30	9
O7–O9	I	-5.48 ± 0.28	0.84 ± 0.09	5	0
O7–O9	V/III	-4.93 ± 0.11	0.86 ± 0.02	1	-3.88 ± 0.36	0.83 ± 0.09	6
WN7	I	-5.86 ± 0.43	0.75 ± 0.08	2	-4.35 ± 0.43	0.75 ± 0.08	2
O6O7If	I	-5.87 ± 0.11	1.03 ± 0.02	1	-4.36 ± 0.17	1.03 ± 0.02	1
O8O9If	I	-7.32 ± 0.10	1.24 ± 0.01	1	-5.81 ± 0.16	1.24 ± 0.01	1
cLBV	I	-7.53 ± 0.10	1.09 ± 0.02	1	-6.03 ± 0.16	1.09 ± 0.02	1

Notes.^a Cluster kinematic distance (Messineo et al. 2008; Reid et al. 2009).^b W33 distance (Immer et al. 2013).^c Luminosity classes are photometrically assigned: for B0–B3 supergiants $M_K < -5.0$ mag, for B0–B3 giants or dwarfs $M_K > -5.0$ mag. For O7–O9 supergiants $M_K < -5.0$ mag, for O7–O9 giants or dwarfs $M_K > -5.0$ mag.^d When Nstar > 1, quoted errors are the standard deviations.

the $^{13}\text{CO } J = 3 - 2$ transition yielded a radial velocity $V_{\text{LSR}} = 50.2 \pm 4.1 \text{ km s}^{-1}$ (Beaumont & Williams 2010). We calculated the distance to N10 using the Galactic rotation curve parameters determined by Reid et al. (2009) from fitting trigonometric parallaxes of star-forming regions, i.e., $R_\odot = 8.4 \pm 0.6 \text{ kpc}$ and $\Theta_0 = 254 \pm 16 \text{ km s}^{-1}$. We find a distance of $4.29^{+0.17}_{-0.19} \text{ kpc}$. This could correspond to the high-velocity component seen in direction of W33; the low-velocity component of W33 (35 km s^{-1}) is also detected along the line of sight of N10 (Dame et al. 2001). Bubble N10 is among the 28% of bubbles currently interacting with a molecular condensation (Deharveng et al. 2010); this clump (hereafter BGPS1865) is located on the western edge of the bubble and has a mass $> 1600 M_\odot$ (Ma et al. 2013, and references therein).

We detected one O4–6 star (star #11, $A_{K_s} = 2.28 \pm 0.03 \text{ mag}$) in the center of the bubble N10, as well as three early type-stars (spectral type OBAF; #12, #15, and #16) and the embedded massive star #14. For a distance of 2.4 kpc, we obtain $M_K = -4.07 \pm 0.1 \text{ mag}$ for #11 (which is a typical value for dwarfs of $29 \pm 3 M_\odot$; Martins & Plez 2006; Ekström et al. 2012); for a kinematic distance of 4.29 kpc this is revised upward to $M_K = -5.08 \pm 0.1 \text{ mag}$ ((super-)giant of $36 \pm 4 M_\odot$; Martins & Plez 2006; Ekström et al. 2012). There is a hint for Si IV in emission at $2.428 \mu\text{m}$ in the spectrum of star #11. Given the uncertainty in reddening and temperature of the remaining four stars, we are unable to determine their luminosities.

8. THE RELATIONSHIP BETWEEN CL1813–178 AND W33

The young massive cluster Cl 1813–178 is projected onto the northwest periphery of W33 and is coincident with SNR G12.72–0.00 (e.g., Brogan et al. 2006; Helfand et al. 2006; Messineo et al. 2008, 2011). Messineo et al. also suggested a possible association of this cluster, with G12.83–0.02 and the pulsar and TeV gamma-ray source PSR J1813–1749/HESS J1813–178 (Brogan et al. 2005).

The cluster contains six spectroscopically detected late O-type stars, 12 early B-type stars, two WN7 stars, and three

transitional objects (the O6O7If star #5, the O8O9If #16, and the cLBV #15 in Messineo et al. 2011). Messineo et al. (2011) estimated a spectrophotometric distance of $3.7 \pm 1.7 \text{ kpc}$, consistent within errors with a stellar kinematic distance of $4.8 \pm 0.2 \text{ kpc}$, which is based on the radial velocity of the RSG member ($V_{\text{LSR}} = 62 \pm 4 \text{ km s}^{-1}$; Messineo et al. 2008) and on the Galactic rotation parameters presented by Reid et al. (2009).

Given our current understanding of stellar evolution, it appears difficult to reconcile the distance of Cl 1813–178 with the new parallactic distance to the W33 complex ($\sim 2.4 \text{ kpc}$). On the basis of the spectral features, luminosity classes can be inferred only for the transitional objects, the two WRs, and two other B0–B3 stars with He I at $2.058 \mu\text{m}$ in emission. While the magnitudes of O7–O9 stars and B0–B3 stars are consistent with both distances of 4.8 and 2.4 kpc, the shorter distance would lead to extremely low luminosities for the three transitional objects, as shown in Table 6. For example, for the O6–O7If star, $M_K = -5.87 \pm 0.11 \text{ mag}$ at 4.8 kpc or $-4.36 \pm 0.17 \text{ mag}$ at 2.4 kpc; the latter value is not compatible with the supergiant class. The spectrum of the O6–O7If star resembles that of star F15 in the Arches cluster (Martins et al. 2008); for F15 we derive $M_K = -5.76 \text{ mag}$ by using the photometry by Figer et al. (2002), the extinction law by Messineo et al. (2005), and a distance of 8.4 kpc. Therefore, the O6–O7If star must be located behind the W33 complex.

Furthermore, for a distance of 2.4 kpc, the cluster members of spectral type O or B would have initial masses below $25 M_\odot$, i.e., below the theoretical and observed lower mass limit for the progenitors of late WN stars (Georgy et al. 2012; Messineo et al. 2011). As a consequence, it would be difficult to understand the presence of the two WN7 (#4, #7 in Messineo et al. 2011), one O8–9If/WN9h, and one O6–O7If star, given the resultant absence of a progenitor population. The average $M_K = -5.86 \text{ mag}$ ($\sigma = 0.43 \text{ mag}$) inferred for the two WN7 cluster members at 4.8 kpc is consistent with that of WN7 stars (-5.38 mag with a $\sigma = 0.38 \text{ mag}$) in Westerlund 1 (Crowther et al. 2006a).

Table 7
Infrared Measurements of the Spectroscopically Detected Stars

ID	2MASS			DENIS			UKIDSS			Flag	GLIMPSE				MSX	WISE		NOMAD
	<i>J</i>	<i>H</i>	<i>K_S</i>	<i>I</i>	<i>J</i>	<i>K_S</i>	<i>J</i>	<i>H</i>	<i>K</i>		[3.6]	[4.5]	[5.8]	[8.0]		<i>W1</i>	<i>W2</i>	
	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)		(mag)	(mag)	(mag)	(mag)		(mag)	(mag)	
1	10.168	8.887	8.024	13.969	10.053	8.059	2	7.247	6.686	6.520	6.192
2	15.364	15.171	12.121	15.252	13.922	13.266	1
3	13.230	11.658	10.913	...	13.203	10.831	13.267	11.592	10.907	1	10.364	10.462	10.058	10.631	10.170	...
4	11.756	10.324	9.653	16.440	11.561	9.668	11.561	10.231	9.500	2	9.008	8.893	8.686	8.493	...	9.050	8.896	...
5	13.018	14.124	13.120	...	13.032	12.640	12.383	1	14.45
6	11.718	10.301	9.466	1
7	9.376	8.508	8.090	12.117	9.457	8.008	9.389	2	7.737	7.764	7.661	7.730	14.07
8	14.913	11.581	9.892	...	14.629	9.734	15.104	11.675	9.805	2	8.588	8.296	7.955	8.126	...	8.627	8.056	...
9	...	13.858	17.322	14.083	12.377	1
10	14.697	12.416	11.185	1
11	...	11.713	10.363	10.249	14.325	11.703	10.278	2	9.401	9.127	8.835
12	13.372	12.081	...	12.907	12.543	12.505	1
13	12.718	11.524	10.826	16.555	12.634	10.698	12.681	11.507	10.894	1
14	12.224	16.918	14.036	12.156	1	10.408	9.677	10.118	9.428	8.128	...
15	13.995	13.189	...	15.927	13.880	12.056	14.006	13.271	13.023	1	17.07
16	13.535	12.944	...	15.439	13.374	...	13.538	12.896	12.633	1	16.85
17	15.411	13.290	12.102	...	14.942	11.975	15.454	13.367	12.175	1	11.305
18	13.457	11.652	10.567	18.408	13.375	10.523	13.526	11.648	10.603	1
19	14.654	13.826	13.053	16.450	14.639	...	14.699	13.984	13.797	1	16.72
20	15.044	14.084	12.918	16.900	14.915	...	15.502	14.881	14.638	1
21	13.531	12.692	11.914	15.770	13.464	11.944	13.538	12.829	12.488	1	17.07
22	15.983	12.664	10.448	10.569	15.421	12.369	10.072	2	8.006	7.283	6.627	5.962	...	7.868	6.766	...
23	10.062	8.741	8.103	14.030	9.997	8.091	9.846	...	7.946	2	7.604	7.493	7.466	16.65
24	12.724	11.505	10.986	15.478	12.838	10.989	12.724	11.500	11.036	1	17.10
25	...	13.023	10.989	10.989	17.218	12.996	10.970	1	9.636	9.516	9.148	9.356	...	9.342	9.077	...
26	10.318	9.128	8.607	13.144	10.164	8.545	10.278	...	8.652	2	8.323	8.362	8.160	8.200	...	8.425	8.419	15.18
27	12.089	10.985	10.546	14.682	11.923	10.514	12.059	10.960	10.592	2	10.232	10.385	9.741	9.343	...	10.251	9.900	16.17
28	...	11.343	10.069	...	14.130	10.280	2	9.197	9.273	8.896	8.929	...	9.232	9.126	...
29	...	13.950	12.888	16.656	14.289	13.069	3	12.040	12.072	11.984
30	14.448	11.369	9.839	...	14.534	9.825	2	8.694	8.764	8.320	8.301	...	9.002	8.648	...
31	...	12.221	9.892	9.827	16.645	12.221	9.796	2	8.197	8.104	7.627	7.768	...	8.452	8.040	...
32	14.969	11.788	14.681	10.123	15.392	11.835	10.081	2
33	...	13.074	10.931	11.053	17.542	13.155	10.921	1	9.415	9.390	8.896	9.114	...	9.651	9.325	...
34	13.342	15.133	13.105	10.772	13.464	12.651	12.363	1	16.69
35	14.043	11.403	1	9.429	9.049	...
36	15.959	13.018	1
37	...	12.807	10.785	16.846	13.052	10.779	1	9.256	9.128	8.682	8.898	...	9.315	8.845	17.20
38	...	12.477	10.530	10.431	...	12.532	10.516	1	9.111	9.032	8.567	8.698	...	9.363	9.007	...
39	12.522	11.127	10.530	16.098	12.530	10.516	12.461	11.216	10.489	3	10.056	10.088	9.952	10.334	10.153	...
40	12.236	11.290	10.841	10.802	12.283	11.310	10.965	2
41	...	12.569	10.368	10.386	17.008	12.560	10.313	2	8.750	8.632	8.171	8.274	...	8.903	8.198	...
42	13.305	10.899	9.761	9.855	2	8.978	9.152	8.785	8.892	...	9.042	9.091	...
43	12.808	10.941	10.124	17.603	12.943	10.083	2	9.557	9.538	9.405	9.470	...	9.595	9.665	...
44	13.968	11.433	10.026	...	14.263	10.288	2	8.799	8.553	8.178	7.981	...	8.606	8.202	...
45	14.392	12.009	10.893	...	14.597	11.000	14.308	12.033	10.879	3	10.193	10.215	9.907	9.945
46	12.919	11.298	10.593	17.500	12.991	10.714	12.757	11.432	10.607	3	10.111	10.191	9.971	9.888	...	10.195	10.227	...
47	11.302	8.275	6.604	...	12.227	7.338	2	5.859	6.023	4.934	4.659	4.173	5.914	4.960	...
48	...	12.102	10.448	15.308	12.232	10.695	3
49	...	12.343	10.054	9.894	16.907	12.242	9.911	2	8.469	8.309	7.934	8.103	...	8.623	8.238	...

Table 7
(Continued)

ID	2MASS			DENIS			UKIDSS			Flag	GLIMPSE				MSX A	WISE		NOMAD R
	<i>J</i>	<i>H</i>	<i>K_S</i>	<i>I</i>	<i>J</i>	<i>K_S</i>	<i>J</i>	<i>H</i>	<i>K</i>		[3.6]	[4.5]	[5.8]	[8.0]		<i>W1</i>	<i>W2</i>	
	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)		(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)
50	...	12.402	9.991	9.864	17.441	12.554	9.971	2	8.185	7.804	7.372	7.380	...	8.345	7.507	...
51	15.406	15.207	8.948	15.483	11.260	9.029	1
52	...	13.189	10.908	11.030	18.233	13.486	11.088	1	9.378	9.157	8.802	8.914	...	9.656	9.153	...
53	...	12.879	10.554	10.341	17.624	12.909	10.466	1	8.842	8.678	8.213	8.258	...	9.192	8.700	...
54	...	11.876	9.650	...	16.250	9.491	16.570	11.990	9.599	2	8.104	7.997	7.471	7.575	...	8.038	7.608	...
55	...	14.413	12.678	18.239	14.487	12.653	1	11.443	11.259	11.026
56	10.359	8.796	8.074	14.426	10.261	8.171	10.272	...	8.390	2	7.647	7.861	7.551	7.584	...	7.708	7.902	18.21
57	...	10.797	8.532	14.561	10.353	...	2	6.713	5.976	5.456	5.432	5.965	6.669	5.790	...
58	15.267	13.008	1
59	11.213	1	8.730	8.409	7.943	7.922
60	10.983	10.173	9.733	12.567	11.113	9.816	10.938	10.119	9.976	2	12.38
61	...	11.646	9.500	9.332	16.123	11.743	9.392	2	7.871	7.812	7.317	7.462	...	8.263	7.906	...
62	15.359	11.039	8.877	1	7.684	7.327	...
63	...	13.134	10.946	10.815	17.622	13.143	10.890	1	9.412	9.277	8.807	9.079	...	9.682	9.328	...
64	15.078	13.038	1
65	...	12.565	10.516	16.647	12.667	10.534	1	9.065	8.912	8.525	8.784	...	9.266	8.899	...
66	15.982	13.826	1
67	15.249	11.288	9.247	...	15.581	9.271	15.312	11.562	9.341	2	7.967	7.485	7.035	7.010	...	8.355	7.534	...
68	10.766	10.654	16.547	12.617	10.705	1	9.434	9.310	8.951	8.993
69	12.400	15.949	13.533	12.488	1	11.679	11.694
70	9.863	8.703	8.271	12.260	9.801	8.243	9.689	...	8.248	2	14.59
71	12.300	9.616	8.313	...	12.316	8.424	2	7.347	7.563	7.112	7.140	...	7.551	7.476	...
72	17.829	13.790	11.754	1	10.010	9.864	...
73	13.698	11.808	10.830	13.851	11.891	11.029	1
74	...	13.951	16.528	14.797	13.974	1
75	15.863	14.412	1
76	9.682	8.697	8.351	11.487	9.656	8.361	...	10.723	8.752	2	8.216	8.268	8.202	7.849	7.762	13.15
77	15.058	12.899	11.936	...	15.116	11.803	15.083	12.965	12.026	1	11.369	11.088	10.888	11.088
78	14.697	12.579	11.684	...	14.524	11.579	14.617	12.601	11.698	1	11.050	10.898	10.814	10.222
79	14.802	12.587	11.665	...	14.527	11.717	14.759	12.634	11.679	1	11.028	10.899	10.928
80	13.517	12.010	10.395	15.174	13.668	10.284	13.497	12.047	10.353	1	9.094	9.046	8.599	8.595
81	15.107	12.377	11.174	...	14.763	11.066	15.158	12.430	11.195	1	10.278	10.134	9.907
82	0
83	10.041	7.668	6.538	16.385	9.916	6.406	2	6.055	6.298	5.556	5.655	...	5.821	5.700	19.11
84	8.914	7.061	6.169	15.116	9.053	6.262	2	5.380	...	5.099	4.903	4.824	5.529	5.257	...
85	9.217	7.221	6.256	15.880	9.132	6.098	2	7.037	6.028	5.355	5.246	4.897	5.642	5.308	...
86	10.410	7.856	6.586	...	10.366	6.539	2	5.917	6.072	5.306	5.223	5.302	5.825	5.388	...
87	13.478	9.616	7.686	...	13.446	7.654	2	6.472	6.093	5.550	5.473	5.406	6.384	5.869	...
88	15.278	12.589	10.172	16.571	15.264	10.091	2	8.484	8.273	7.760	7.800	...	8.891	8.200	...
89	11.464	11.362	...	14.940	11.465	1
90	...	12.981	11.020	17.045	13.036	11.008	1	9.590	9.522	9.172
91	14.749	12.311	11.216	...	14.528	11.346	14.767	12.317	11.228	1	10.451	10.450	10.422	10.482	10.522	...
92	13.001	11.266	10.557	17.439	12.942	10.519	13.328	11.500	10.690	1	10.082	10.117	10.129	10.333	...	10.022	9.973	19.88
93	17.757	13.992	12.169	1

Notes. The Flag column indicates which *JHK* was adopted in the paper: Flag = 0 source missing; Flag = 1 UKIDSS psf-fitting magnitudes; Flag = 2 2MASS magnitudes; Flag = 3 UKIDSS DR6 release. Stars #59 and #60 are blended even at the UKIDSS resolution, listed measurements should be taken as upper limits. Identification numbers are taken from Tables 1 and 3.

While the SN that gave rise to the remnant G12.72–0.00 may have occurred in Cl 1813–178 (given the precise superposition), a physical association with SNR G12.82–0.02 and associated pulsar PSR J1813–1749 appears doubtful. Specifically, a comparison of the significant column density to the pulsar and SNR to the less extreme extinction inferred for cluster members led Halpern et al. (2012) to conclude that SNR G12.82–0.02 likely lies beyond both the cluster and the W33 complex ($d \sim 5–12$ kpc).

Finally, it is of interest that the distance estimate for the W33 complex was determined from parallax measurement of masers. If it were to be observed in the future, when such emission had ceased, it would be difficult to recognize it as a complex of discrete sources at the same distance and to further distinguish the distinct stellar population of Cl1813–178 that is projected on the edge of W33 (see Figure 7).

9. SUMMARY

We performed a near-IR spectroscopic survey for massive stars, encompassing both W33 and the nearby mid-IR bubble/stellar cluster N10/BDS2003-115 to study their star formation history.

1. We detected a total of 14 new early-type (OB and WR) stars and a further nine stars with spectra consistent with spectral types earlier than F. A large population of giants with spectral types G–M were uncovered, but no cool supergiants associated with W33 were identified.
2. Following Clark et al. (2009b), the lack of RSGs precludes substantive star formation activity with W33 ≥ 6 Myr, while the detected stellar population appears broadly consistent with an age of $\sim 2–4$ Myr.
3. The complex contains protostars (most notably the embedded protocluster W33 Main and the high-mass protostar W33A), massive evolved stars, and clear marks of sequential star formation and feedback. Star formation within W33 has not led to the formation of rich dense clusters, and the size of W33 (radius ≈ 5 pc) is typical for loose associations (Pfalzner 2009). When the GMC is exhausted and star formation has ceased, W33 will most likely resemble a loose, non-coeval stellar association similar to (but less massive than) Cyg OB2 (e.g., Negueruela et al. 2008).
4. Given the sparse nature of individual stellar “aggregates” and the limitations of the current data, we cannot infer integrated masses for the young populations within W33. W33 is probable less massive than other massive star-forming complexes of the Milky Way with known evolved stars, such as W43 (e.g., Blum et al. 1999; Chen et al. 2013), W51 (Clark et al. 2009b), Carina (e.g., Preibisch et al. 2011), and G305 (Davies et al. 2012), as suggested by their respective integrated radio and IR luminosities (e.g., Immer et al. 2013; Conti & Crowther 2004).
5. The greater distance to the nearby young massive stellar aggregate Cl 1813–178 precludes a physical association with W33. The late O-type and B-type members of the cluster support a distinct older population than that observed in W33. Considering the extinction of Cl 1813–178 ($A_V = 9.1$ mag), optical spectroscopy would yield precise spectral types and direct luminosity

determinations for its constituent stars (e.g., Negueruela et al. 2010).

6. Given the distances to W33 and to Cl 1813–178, an association with the energetic pulsar PSR J1813–1749 appears doubtful.

This work was partially funded by the ERC Advanced Investigator Grant GLOSTAR (247078). F.N. acknowledges funding from the Spanish Government Ministerio de Economía y Competitividad (MINECO) through grants AYA2010-21697-C05-01, FIS2012-39162-C06-01 and ESP2013-47809-C3-1 R. This publication makes use of data products from the Two Micron All Sky Survey, which is a joint project of the University of Massachusetts and the Infrared Processing and Analysis Center/California Institute of Technology, funded by the National Aeronautics and Space Administration and the National Science Foundation. This work is based on observations made with the *Spitzer Space Telescope*, which is operated by the Jet Propulsion Laboratory, California Institute of Technology, under a contract with NASA. DENIS is a joint effort of several institutes, mostly located in Europe. It has been supported mainly by the French Institut National des Sciences de l’Univers, CNRS, and French Education Ministry, the European Southern Observatory, the State of Baden-Wuerttemberg, and the European Commission under networks of the SCIENCE and Human Capital and Mobility programs, the Landessternwarte, Heidelberg, and Institut d’Astrophysique de Paris. This research made use of data products from the *Midcourse Space Experiment*, the processing of which was funded by the Ballistic Missile Defence Organization with additional support from the NASA office of Space Science. This publication makes use of data products from *WISE*, which is a joint project of the University of California, Los Angeles, and the Jet Propulsion Laboratory/California Institute of Technology, funded by the National Aeronautics and Space Administration. This work is based in part on data obtained as part of the UKIRT Infrared Deep Sky Survey. This research made use of Montage, funded by the National Aeronautics and Space Administration’s Earth Science Technology Office, Computational Technologies Project, under Cooperative Agreement Number NCC5-626 between NASA and the California Institute of Technology. The code is maintained by the NASA/IPAC Infrared Science Archive. This research has made use of the VizieR catalogue access tool, CDS, Strasbourg, France. This research has made use of the SIMBAD database, operated at CDS, Strasbourg, France. This research has made use of NASA’s Astrophysics Data System Bibliographic Services. A special thanks goes to the great support offered by the European Southern Observatory. M.M. thanks Jos de Bruine and Timo Prusti for useful discussions and support while at ESA. We thank the referee Dr. Philip Dufton for his careful reading.

APPENDIX FINDING CHARTS

Figure 9 displays charts for the spectroscopically observed stars. For a few stars that are not easily identifiable in this figure additional SINFONI charts are provided in Figure 10.

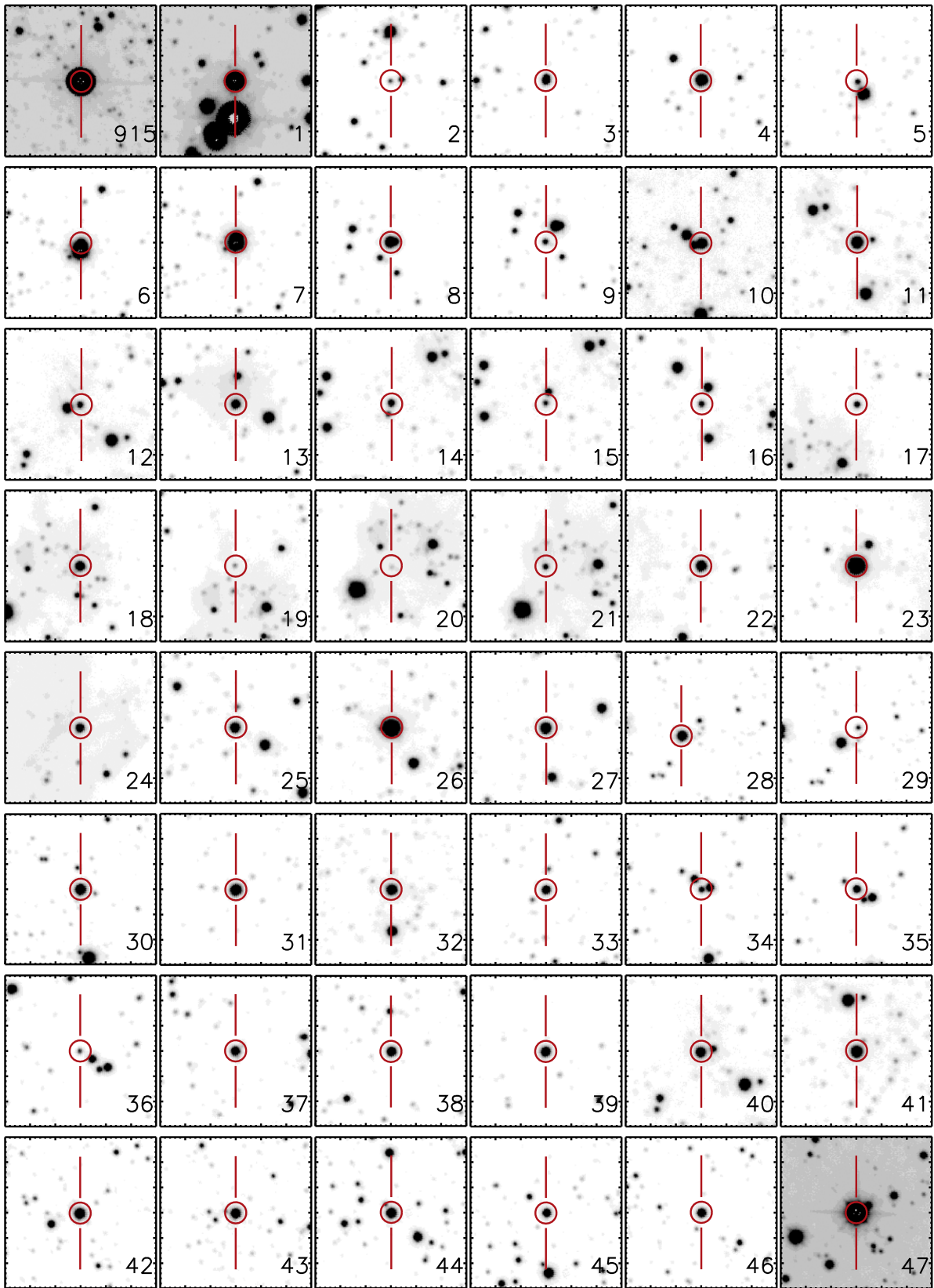


Figure 9. UKIDSS *K*-band charts (0.5×0.5) of stars spectroscopically observed. North is up and east to the left.

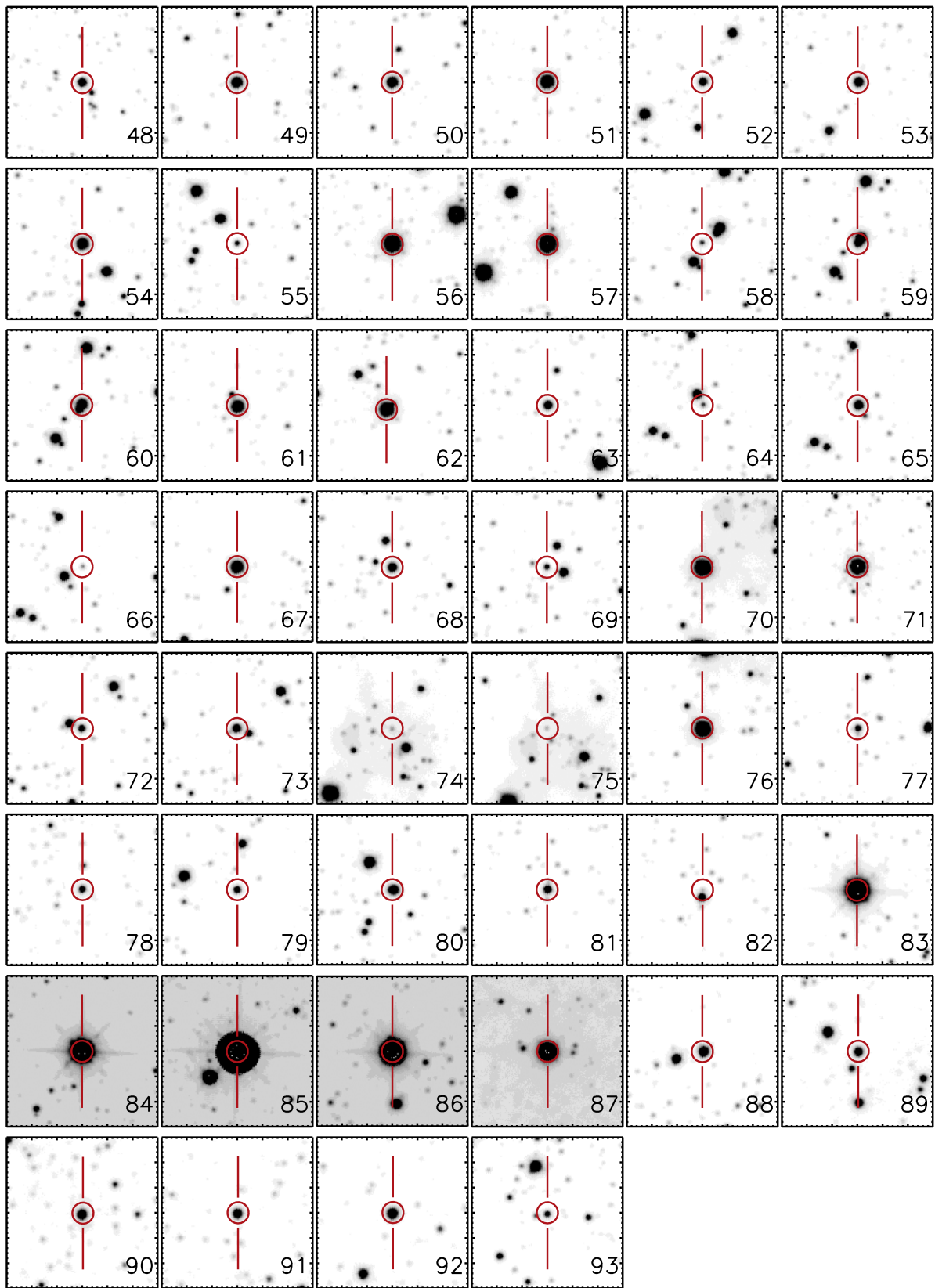


Figure 9. (Continued.)

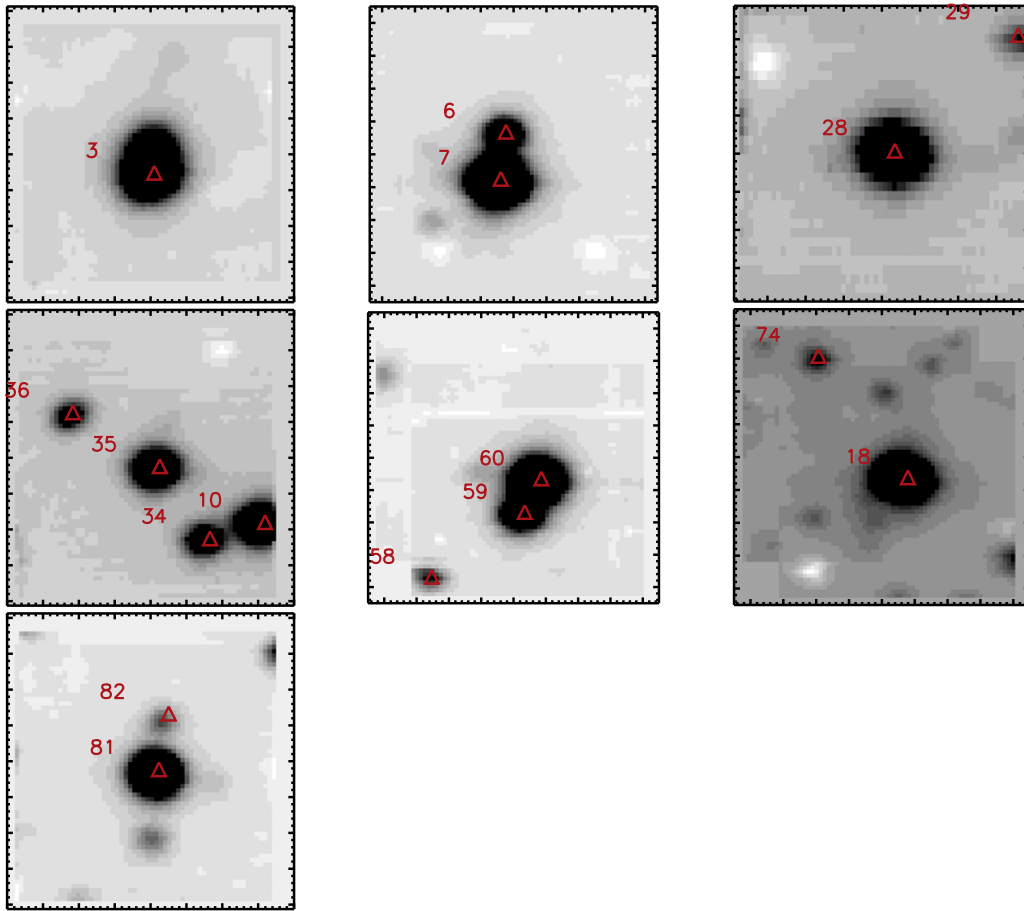


Figure 10. Additional charts for sources not easily identifiable in Figure 9. The displayed images are obtained by averaging the SINFONI data-cubes in wavelength (K grating, $8'' \times 8''$).

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